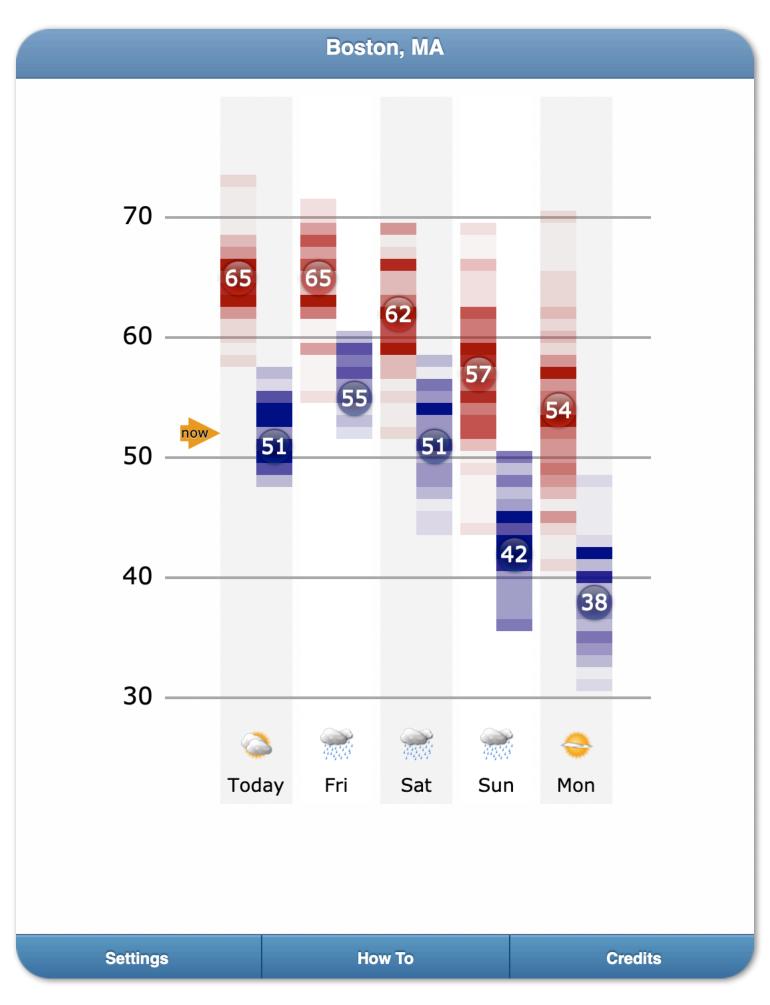
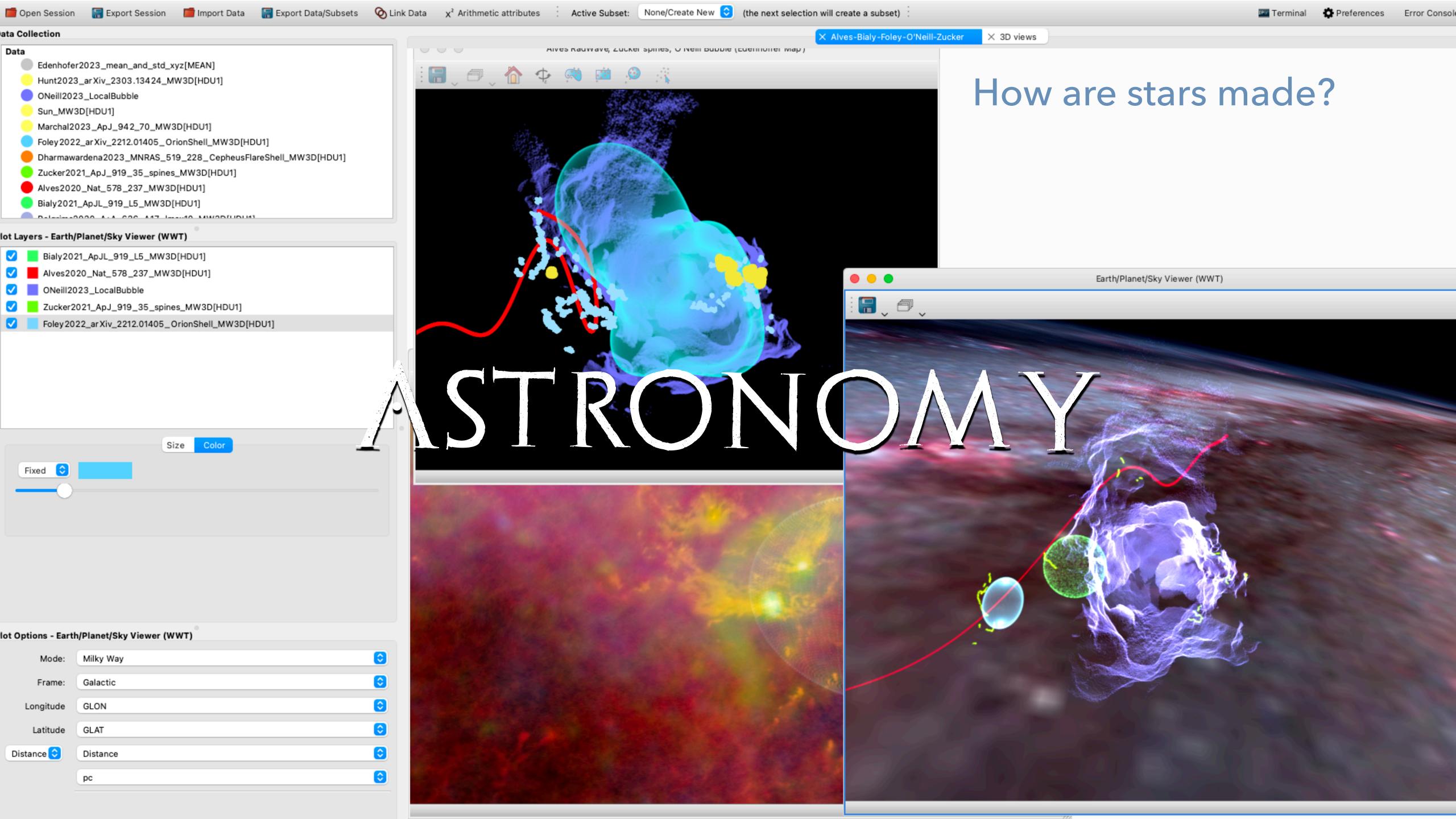


How are weather forecasts made?



today's forecast from "takeasweater.com," part of The Prediction Project (predictionx.org).





2502 P. S. Li, C. F. McKee and R. I. Klein

Table 1. Glossary

eter ^a	Definition	Reference
	Magnitude of initial magnetic field: 31.6 μG (strong), 3.16 μG (weak)	Section 4.2
	Magnitude of mass-weighted line-of-sight field in central Gaussian beam	Equation (1)
z'	Average magnitude of field in central τr_{beam}^2 of clump mid-plane	Section 7
,	Magnitude of average axial field in central τv_{beam}^2 of clump mid-plane	Section 7
	Total mass-weighted field, $(B_{\log_x x}^2 + B_{\log_x y}^2 + B_{\log_x z}^2)^{1/2}$	Equation (22)
	Magnitude of the field averaged over clump volume, $\langle B \rangle_V$	Section 8.1
	RMS field in clump volume, $(\langle B^2 \rangle_V)^{1/2}$	Section 6.2
	Mass-weighted isothermal sound speed	Equation (A5)
	$A/\ell^2 = 1(\text{box}), \pi/4 \text{ (sphere)}$	Equation (A3)
	$V/\ell^3 = 1(\text{box}), \pi/6 \text{ (sphere)}$	Equation (A4)
	$\sigma_{1D}^2/\sigma_{nt}^2 = 1 + 3/\mathcal{M}^2$; $\rightarrow 1$ for highly supersonic flow	Equation (A5)
	Numerical coefficient in M_{Φ} ; adopt $c_{\Phi} = 1/(2\pi)$	Equation (A1)
	Magnetic critical mass	Equation (A1)
	Alternate form of magnetic critical mass, M_{Φ}/μ_{Φ}^2	Equation (A10)
	Sonic Mach number, $v_{\rm rms}/c_{\rm s}$	Section 4.1
	Alfvén Mach number, $v_{\rm rms}/v_{\rm A}$	Section 4.1
	Mean density of H nuclei	Equation (3)
	Mean density of H nuclei in central Gaussian beam	Equation (21)
	Column density of H nuclei in central Gaussian beam	Equation (21)
	Mass-weighted Alfvén velocity, $B_{\text{rms},V}/(4\pi\bar{\rho})^{1/2}$	Section 4.1
	Mass-weighted 3D non-thermal velocity dispersion, $\sqrt{3}\sigma_{\rm nt}$	Section 4.1
_	Average y of x in 100-clump sample; also averaged over time	Section 5
	Expone f density scaling of magnetic field, $B \propto \bar{n}_{\rm H}^{\alpha}$	Equation (3)
	Plasm $= 8\pi\bar{\rho}c_s^2/B_{\text{rms},V}^2$	Equation (11)
1	Vo (nized) mass-to-flux ratio, $M/M_{\Phi}=2\pi G^{1/2}\Sigma/B$	Equation (A2)
	al-beam line-of-sight mass-to-flux ratio, $2\pi G^{1/2} \Sigma_{\rm beam}/B_{\rm los}$	Equation (28)
	$\langle \Phi \rangle^{-1}$ = harmonic mean mass/flux ratio = average flux/mass ratio	Equation (29)
	N s-weighted total 1D velocity dispersion, $(c_s^2 + \sigma_{nt}^2)^{1/2}$	Equation (A5)
	N s-weighted 1D non-thermal velocity dispersion	Equation (A5)
	See density in central Gaussian beam	Section 6

Notes. ^aSubscripts: beam = central Gaussian beam, h = harmonic mean, H = hydrogen nuclei, los = line of sight, M = mass, med = median, t = time-averaged, V = volume-averaged, $\mu = \text{microgauss}$.

observations by Falgarone et al. (2008). The OH data comprises 54 sources, 20 of which are detected at more than 2σ ; the CN data is from 14 sources, 8 of which are detected at more than 2σ . Crutcher et al. (2010) used a Bayesian analysis to self-consistently include upper limits on the field strengths with a statistical significance less than 2σ

The Zeeman signal is proportional both to the integral of the field along the line of sight and to the column density of the molecule being observed. Let the line of sight be along the z direction. We assume that the molecule has a fixed abundance, so that the column density of the molecule is proportional to the mass column density. Then the magnitude of the Zeeman field measured by a given telescope beam is

$$B_{\text{los}} = \frac{1}{\bar{\rho}V} \left| \int_{\text{beam}} \rho B_z dV \right| = \frac{1}{M_{\text{beam}}} \left| \int_{\text{beam}} B_z dM \right|, \tag{1}$$

where M_{beam} is the mass inside the beam and B_z is the component of the magnetic field along the line of sight. (Definitions are collected together in Table 1.) Note that B_z is a component of a vector and can be positive or negative, whereas values of the field that are not explicitly components of a vector, such as B_{los} , are field magnitudes. If the molecule being observed is excited only above some critical density, then M_{beam} is the mass in the beam above that density. It is important to note that if the field is tangled, then B_{los} can be less than

the average magnitude of B_z . Of course, what is of greater interest is the value of the field in the absence of projection effects,

$$\mathbf{B}_{\text{tot}} = \frac{1}{\mathbf{M}_{\text{beam}}} \int_{\text{beam}} \mathbf{B} \, d\mathbf{M},$$
 (2)

and this is the focus of the Crutcher et al. (2010) analysis.

Crutcher et al. (2010) carried out a Bayesian analysis to infer the distribution of $B_{\rm tot}$ as a function of the mean density of hydrogen nuclei, $\bar{n}_{\rm H}$, which extends over the range $10~{\rm cm}^{-3}\lesssim \bar{n}_{\rm H}\lesssim 10^7~{\rm cm}^{-3}$ in their sample. They considered several different distributions and found that the best fit to their data is a uniform distribution of $B_{\rm tot}$ from $fB_{\rm tot,max}$ to $B_{\rm tot,max}$ with

$$B_{\text{tot, max}} = \begin{cases} B_0 & (\bar{n}_{\text{H}} < n_0), \\ B_0 \left(\frac{\bar{n}_{\text{H}}}{n_0}\right)^{\alpha} & (\bar{n}_{\text{H}} > n_0). \end{cases}$$
 (3)

They found $B_0 = 10 \,\mu\text{G}$, $n_0 = 300 \,\text{cm}^{-3}$, $\alpha = 0.65$, and $f = 0.03 \simeq 0$.

The density n_0 dividing the constant density regime from the power-law one can be understood as the critical density, $n_{\rm H,crit}$, at which the thermal pressure in the clump balances the turbulent pressure of the ambient gas,

$$P_{\text{turb}} = \bar{\rho}\sigma_{\text{nt}}^2 = x_{\text{tot}} n_{\text{H, crit}} k_{\text{B}} T, \tag{4}$$

where $\bar{\rho}$ is the mean density, $x_{\rm tot} \equiv n_{\rm tot}/\bar{n}_{\rm H}$ is the total number density relative to hydrogen, and $k_{\rm B}$ is the Boltzmann constant. Since the magnetic field is about the same inside and outside the clump,

"STARFORGE" simulations

MNRAS **452**, 2500–2527 (2015)

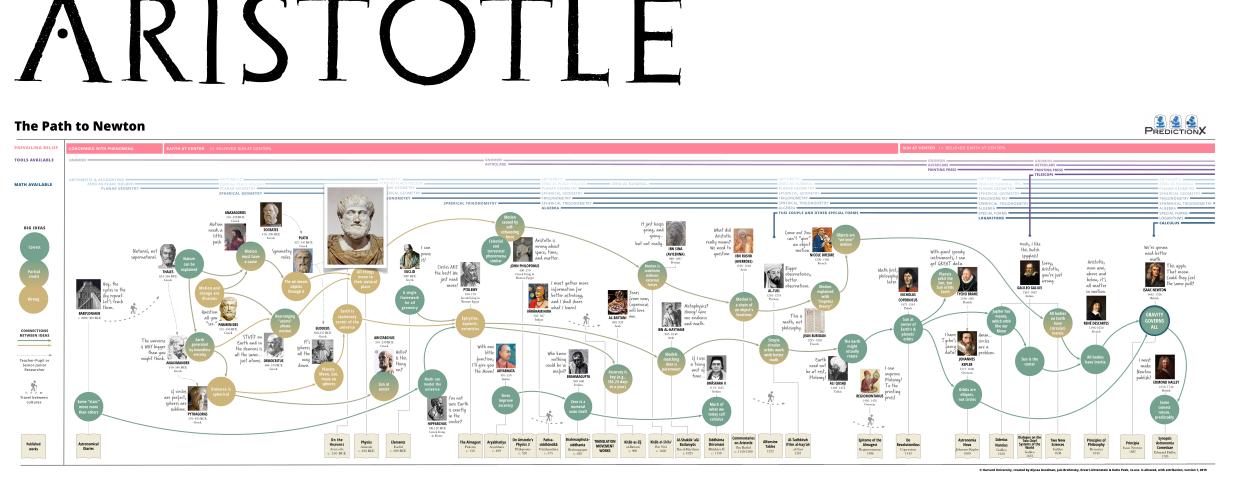
Documentation of the MHD functionality of the "Orion" simulation code, from: Li, Pak Shing, Christopher F. McKee, and Richard I. Klein. 2015. "Magnetized Interstellar Molecular Clouds – I. Comparison between Simulations and Zeeman Observations." Monthly Notices of the Royal Astronomical Society 452 (3): 2500–2527. https://doi.org/10.1093/mnras/stv1437.

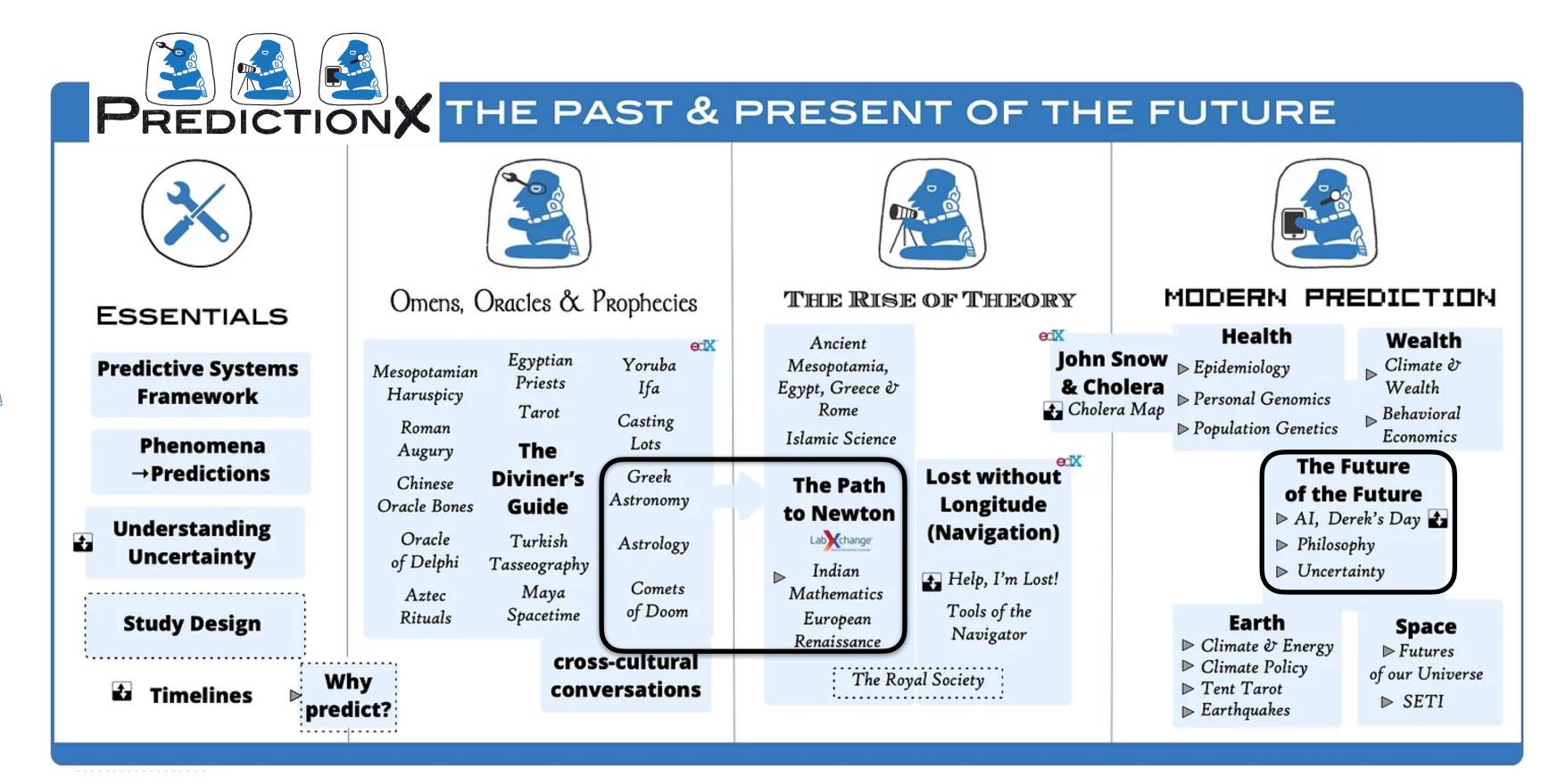
The Prediction Project

The Past and Present of the Future











Phenomenon	Observation	Data	Rule	Theory	Explanation	Prediction
Phenomenon	Observation	Data		 ⑤ OpenAI		Prediction

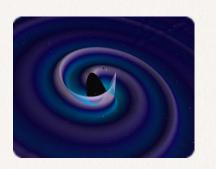
Opinion

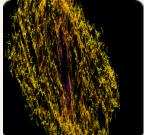
The Aem Aniverse

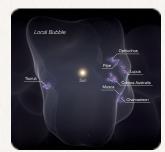
MEMPHIS, SUNDAY OCTOBER 23, 2022

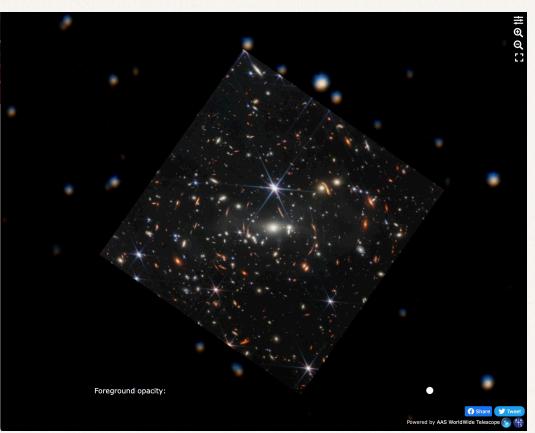
WHAT DO EXPENSIVE NEW TELESCOPES DO FOR HUMANITY TODAY?

Are mega-projects like ALMA, LIGO, JWST, and Gaia worth the billions?



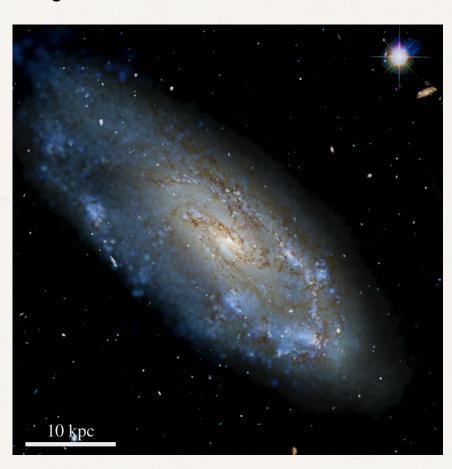


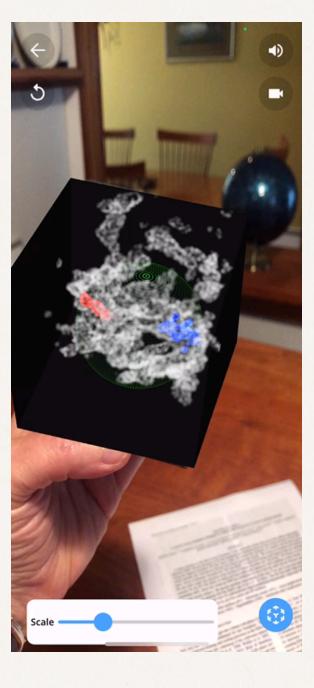




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New galaxies in-silico, the early Universe without physics, and new stars forming in your hand.

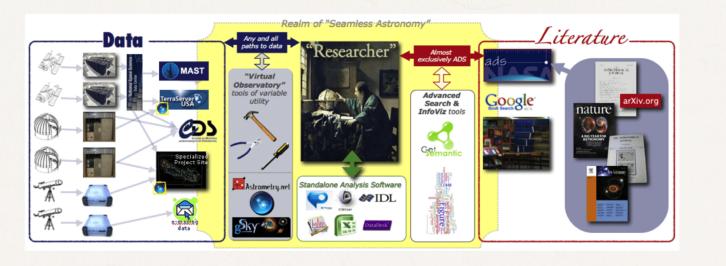




IS ASTROPHYSICS BEING (RE)ORGANIZED?

Lone stargazers are a rarer and rarer breed in professional astronomy. Teams and data scientists seem the way of the future, and tools that talk to each other are essential.





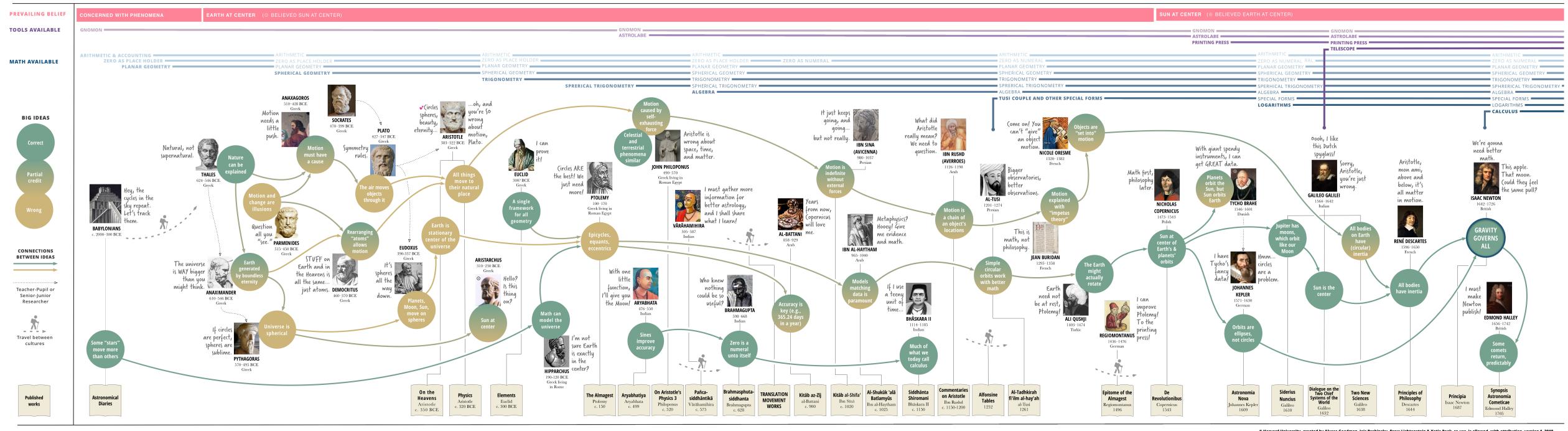
ARE COMPUTERS THE NEW TELESCOPES?

The Path to Newton



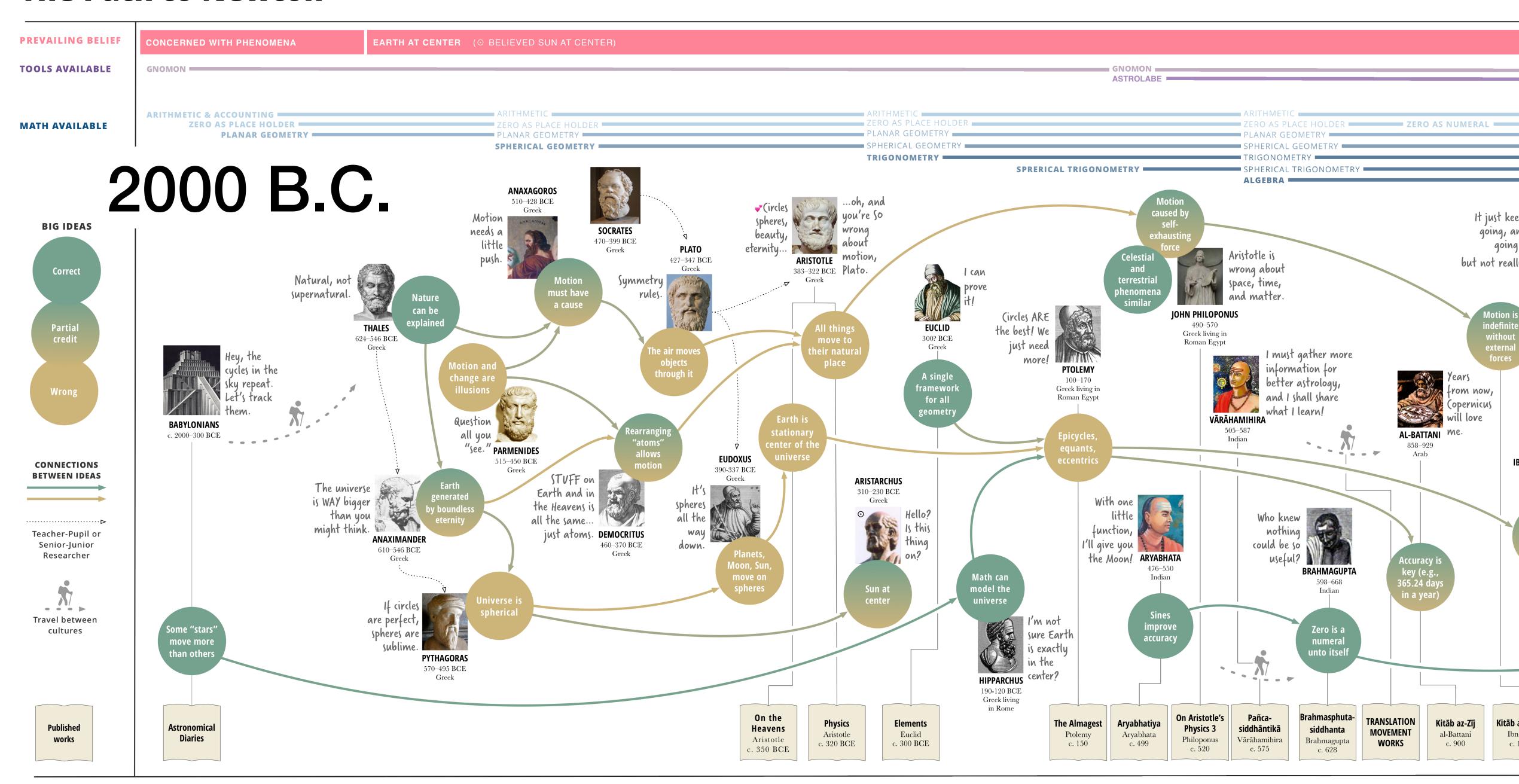
The Path to Newton

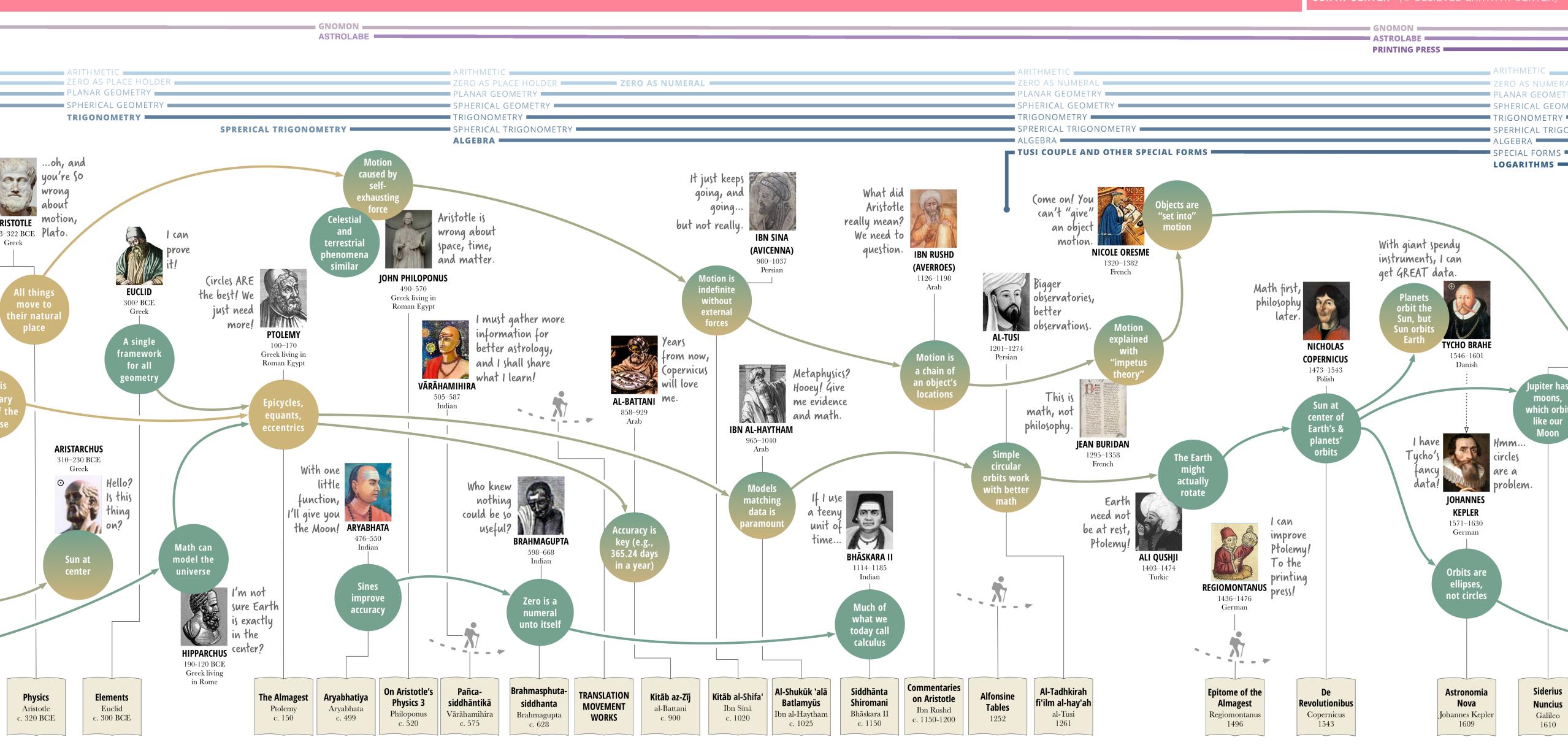




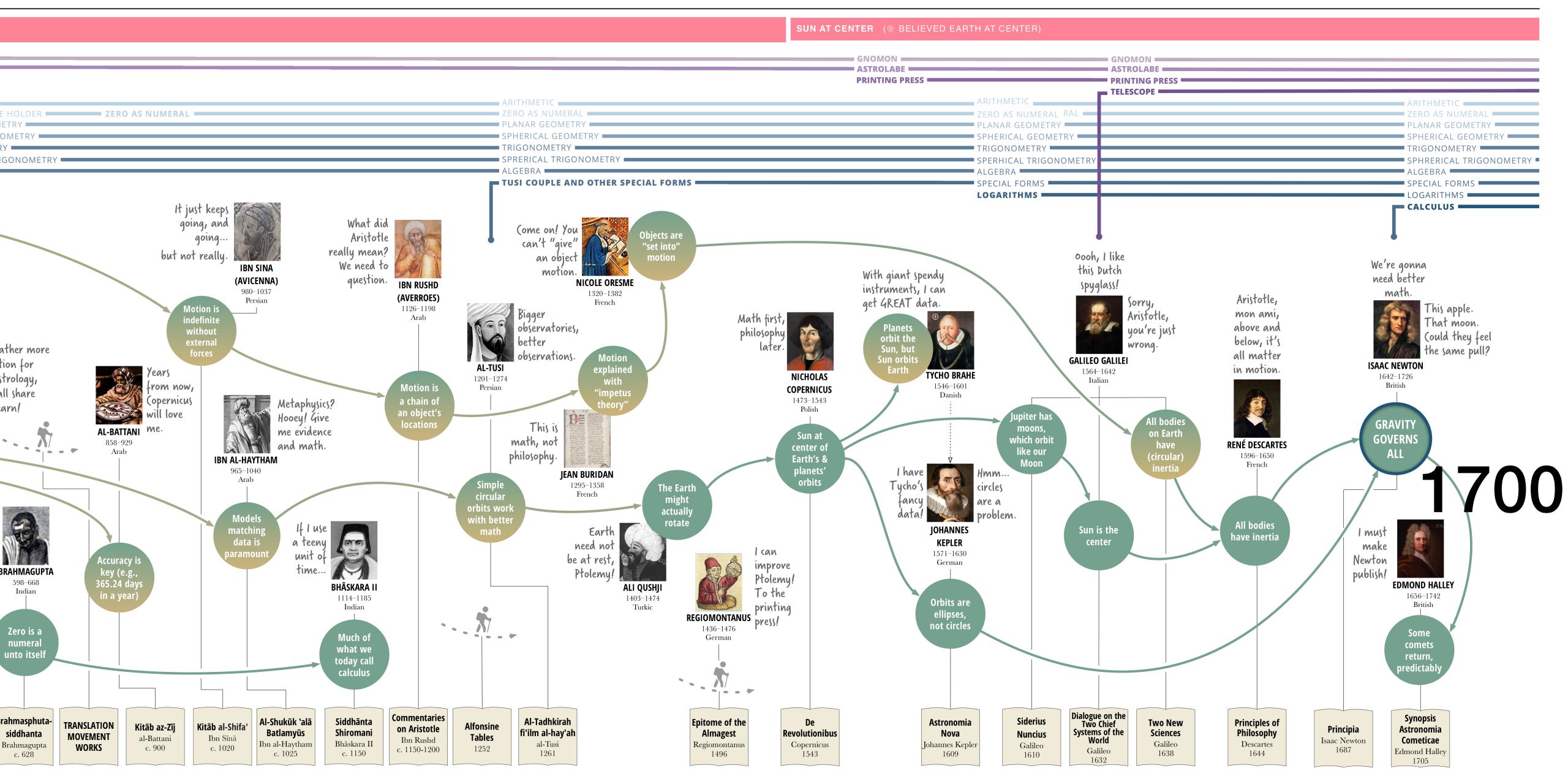
demo: path-to.org or Lab change™

The Path to Newton











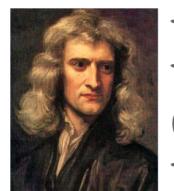
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CALILEI



Aristotle, mon ami, above and below, it's all matter We're gonna need better math.



This apple.
That moon.
(ould they feel the same pull?

"ARE COMPUTERS THE NEW TELESCOPES?"

ELECTRONIC COMPUTERS

NUMERICAL SIMULATION

BAYESIAN STATISTICS

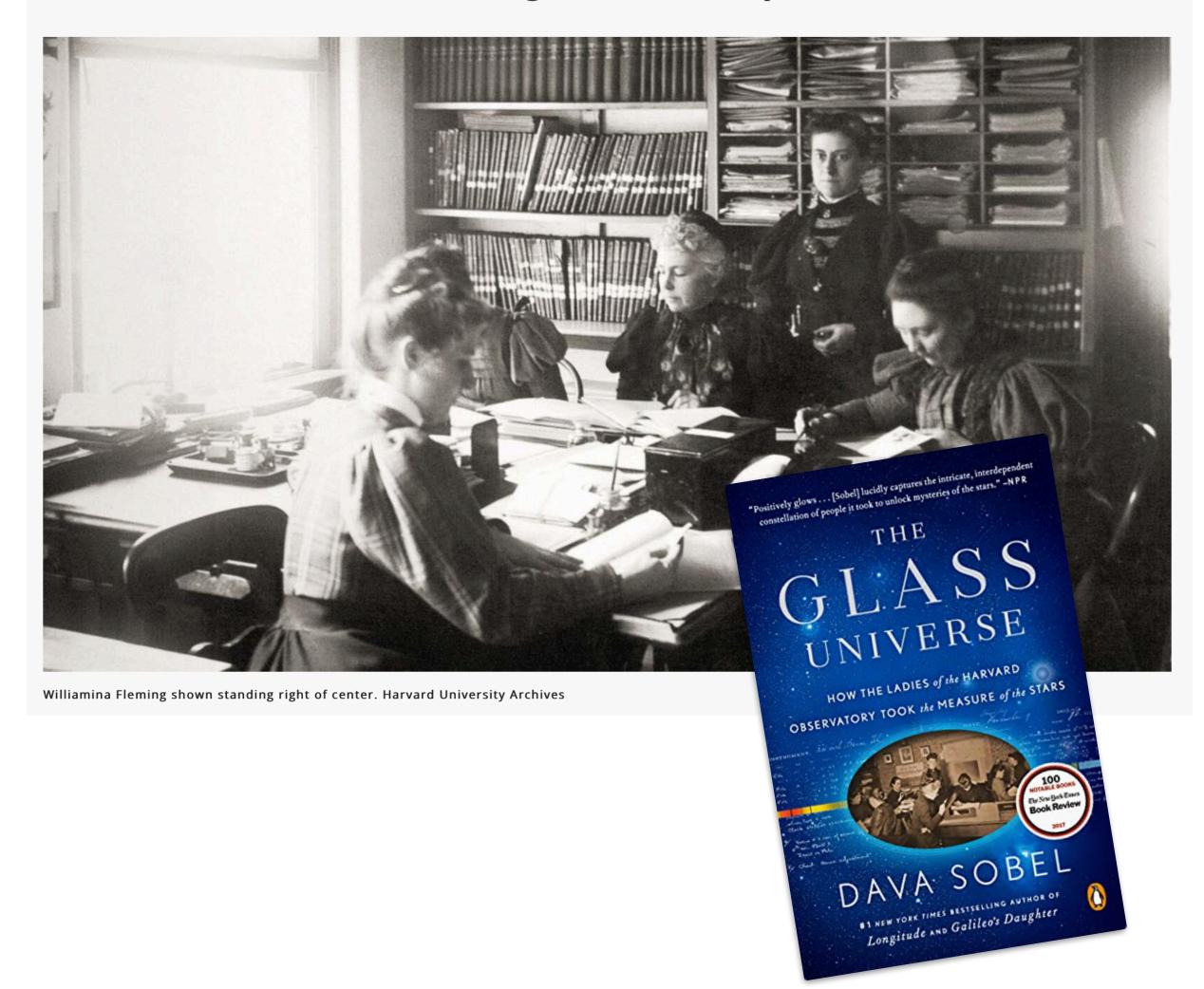
INTERACTIVE DATA VISUALIZATION

AI/MACHINE LEARNING —

DATA SCIENCE IN ASTRONOMY BEGAN — With HUMAN COMPUTERS AT HARVARD

HOME / ASTRONOMICAL PHOTOGRAPHIC GLASS PLATE COLLECTION /

Women at the Harvard College Observatory







the first white dwarf in 1910. She became the first woman with an officially appointed job at Harvard as Curator of Astronomical

Photographs in 1898.



Antonia Maury

First to detect and calculate the orbit of a spectroscopic binary, which was published in 1933. She worked intermittently with Harvard College Observatory before returning in 1918 as an adjunct professor until her retirement in 1948.



Henrietta Swan Leavitt

Found period-luminosity relation for Cepheid variable stars in 1912, providing astronomers with the first "standard candle" with which to measure the distance to faraway galaxies.



Annie Jump Cannon

Developed the Harvard Classification Scheme in 1922, the stellar classification system still in use today. In 1938, two years before retirement, she finally obtained a regular appointment from Harvard as William C. Bond Astronomer.



Cecilia Payne-Gaposchkin

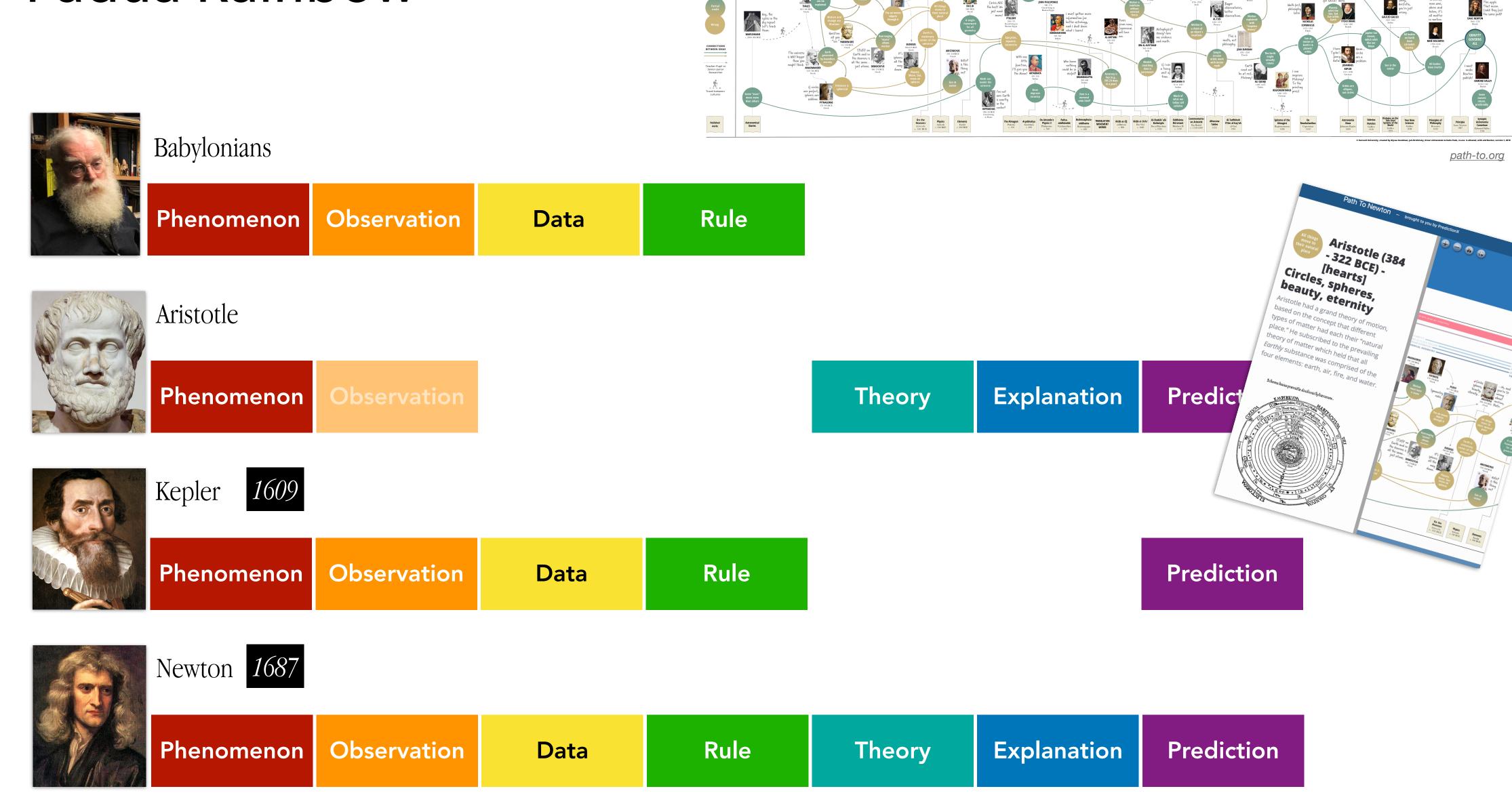
Determined the chemical composition of stars, particularly that hydrogen and helium are the most abundant elements. She was the first person awarded a PhD in Astronomy at Harvard, and in 1956 she became Harvard's first regularly appointed female professor and the first female department chair.



1960

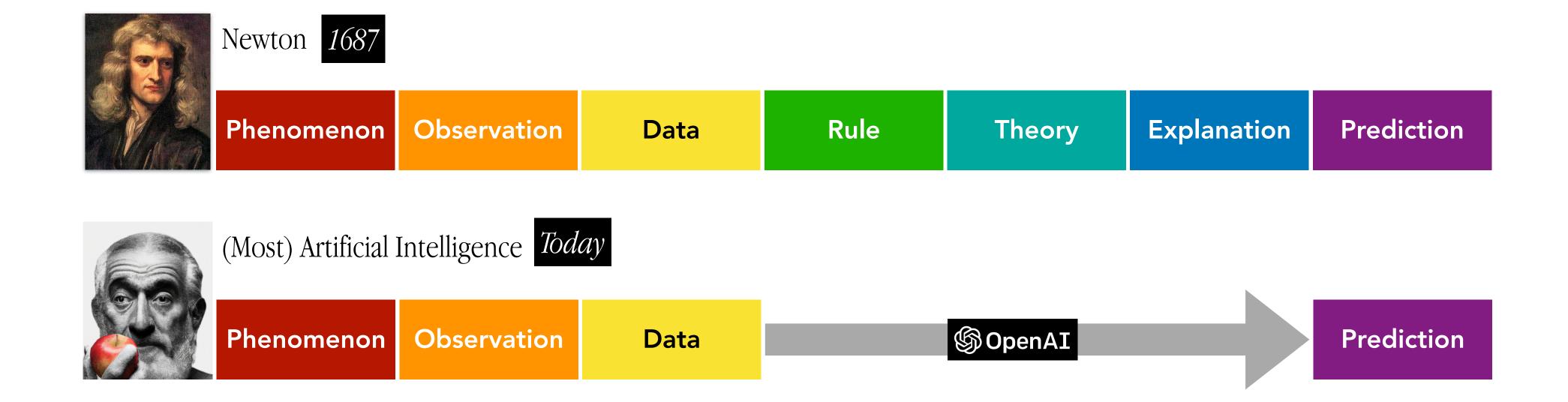
1880

The "Padua Rainbow"

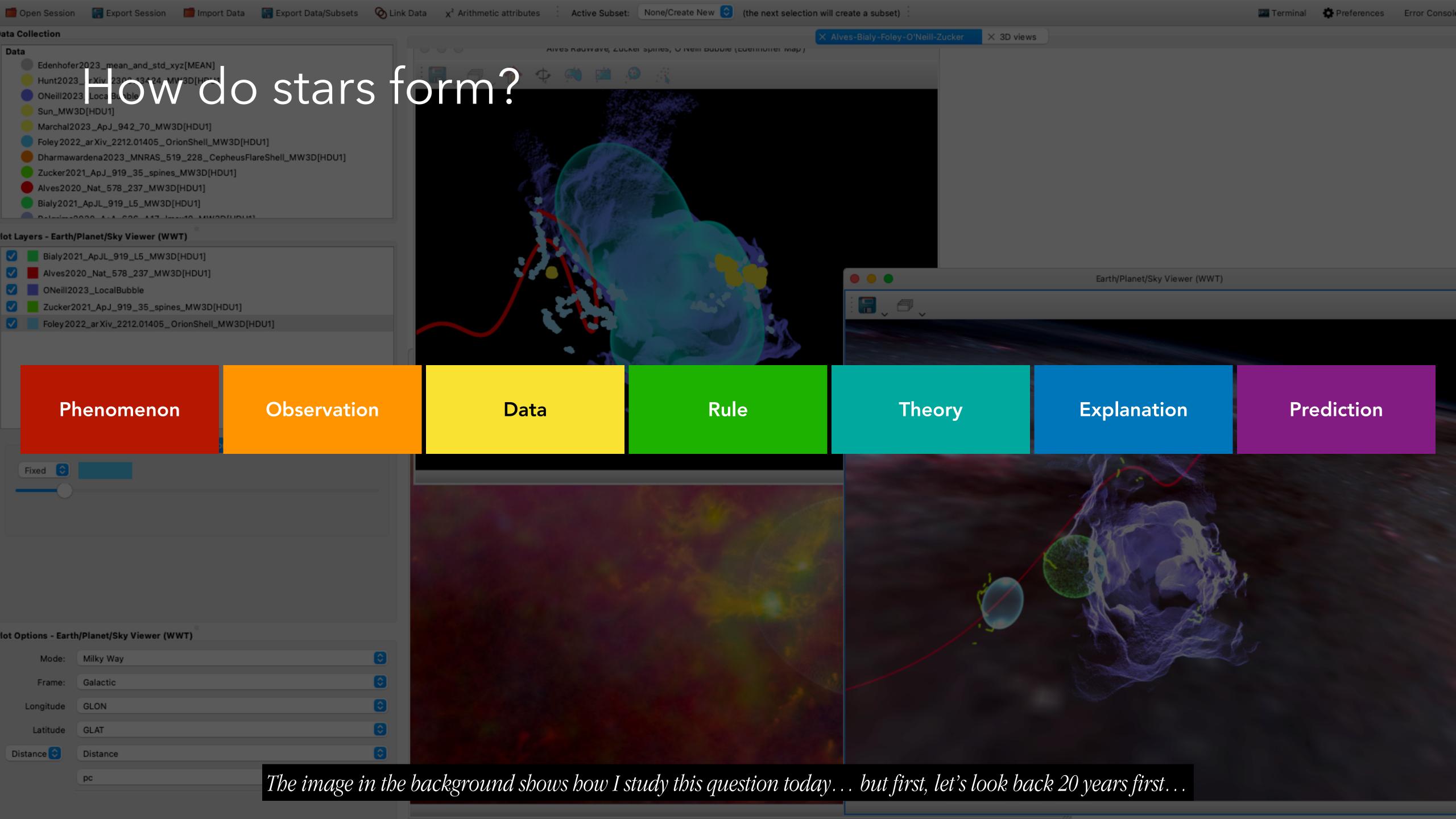


The Path to Newton

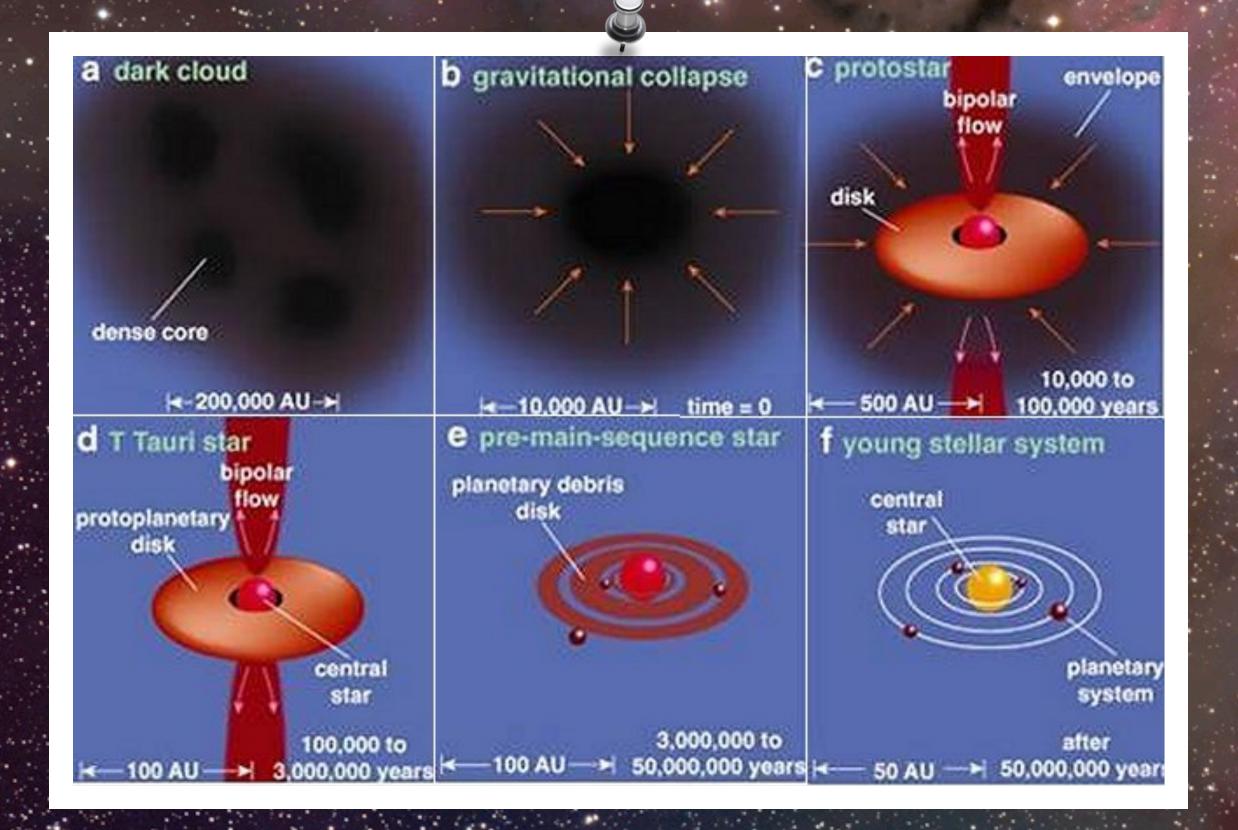
PREDICTIONX

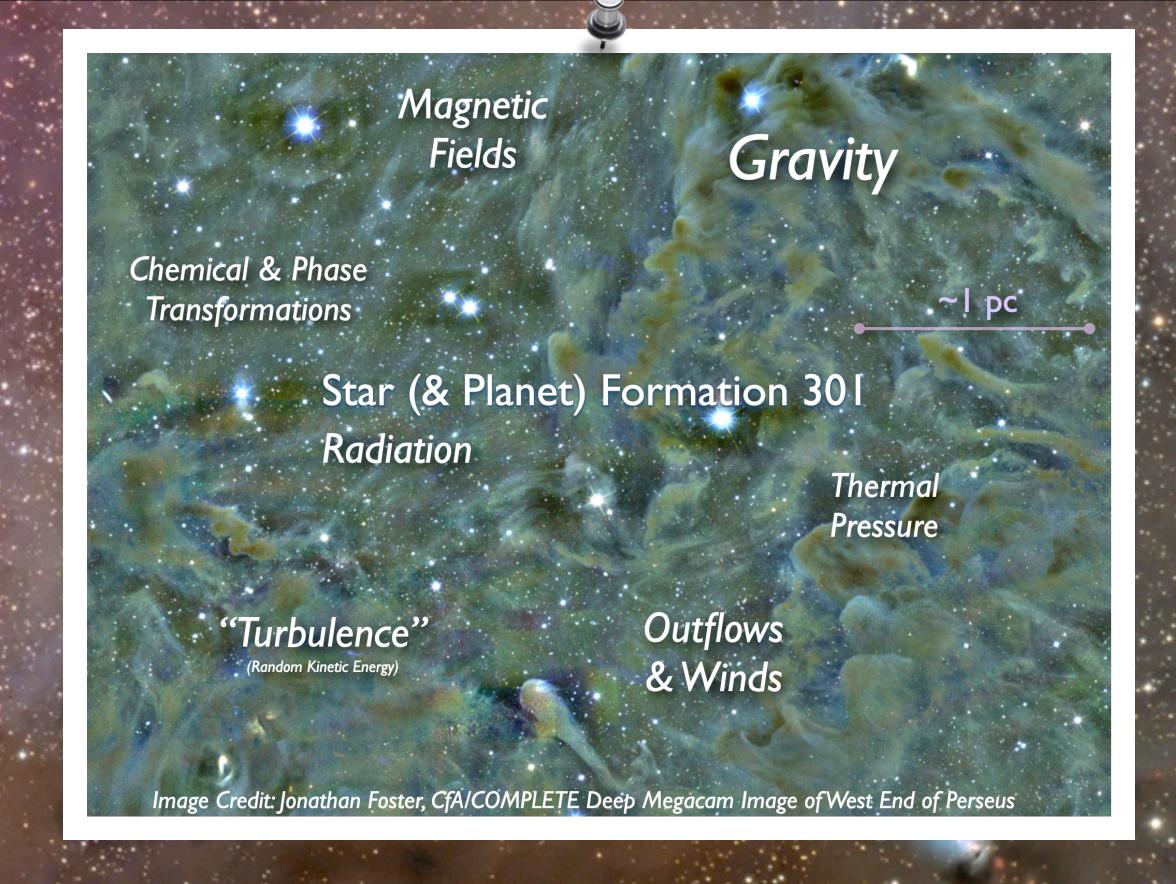


We will come back to this later...



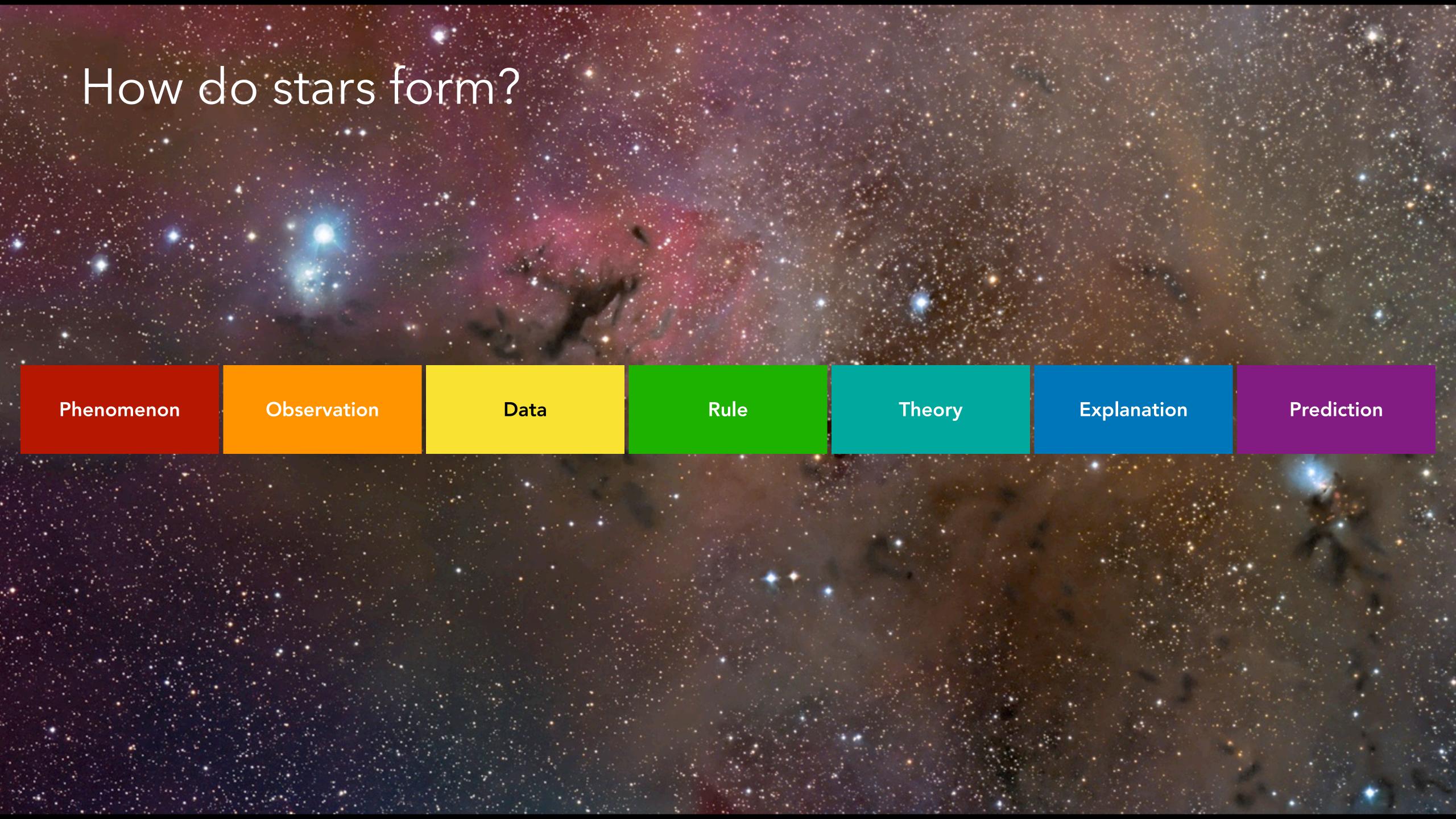
How do stars form?





Perseus

www.flickr.com/photos/66496709@N00/6791649829/in/photostream © Adam Block/Science Photo Library



Phenomenor

Looks like there are stars and clouds of darkness in the Sky.

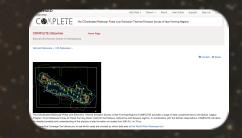
Observation

Let's observe them every way we know how.



Data

And keep track of everything we find.



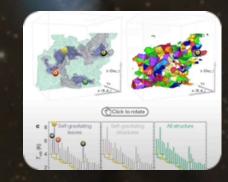
Rule

Bright nebulae are often near dark cloudy patches.



Theory

Stars form from collapsing gas in the clouds.



Explanation

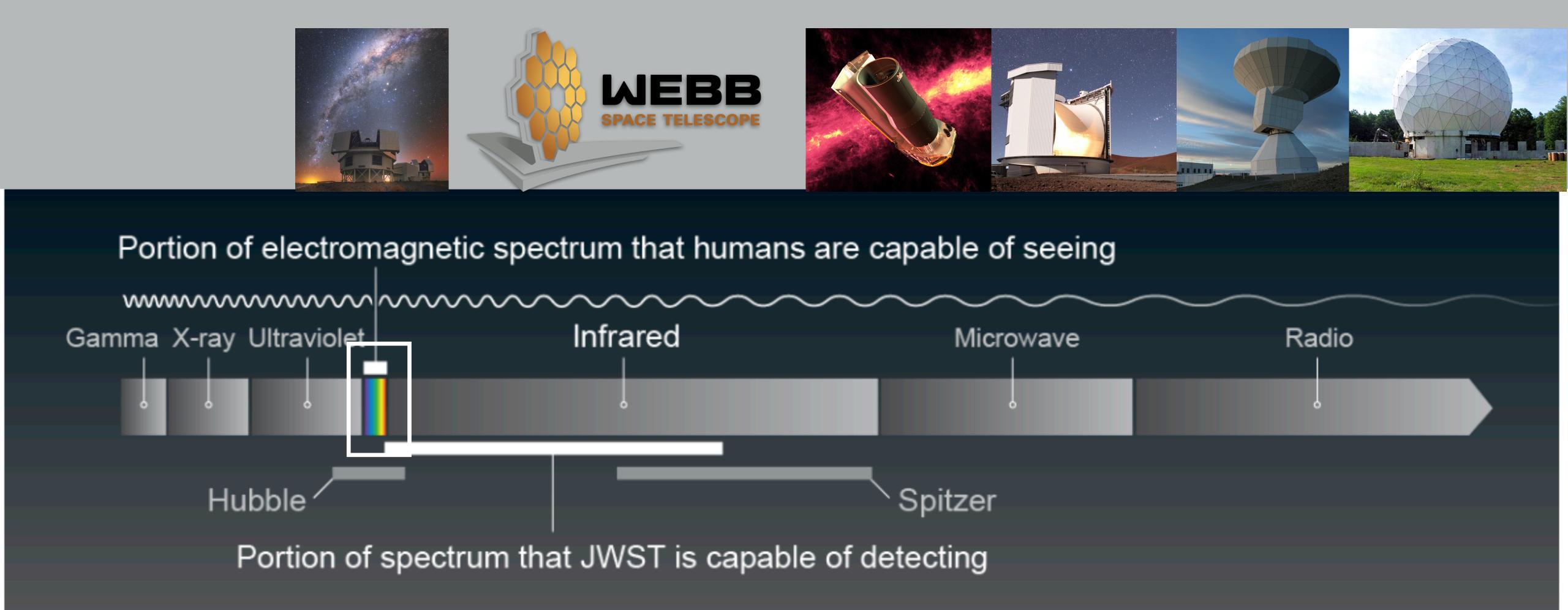
Bright nebulae are caused by new stars lighting up clouds.

Constitution University
Constitution From the Constitution Cons

Prediction

Dark clouds form new stars, even if we can't see them.











Want to see in the INFRARED like JWST can?

Watch the demo read the or X me to start playing right now!

This mini data story is brought to you by NASA's SciAct CosmicDS program and AAS WorldWide Telescope.





















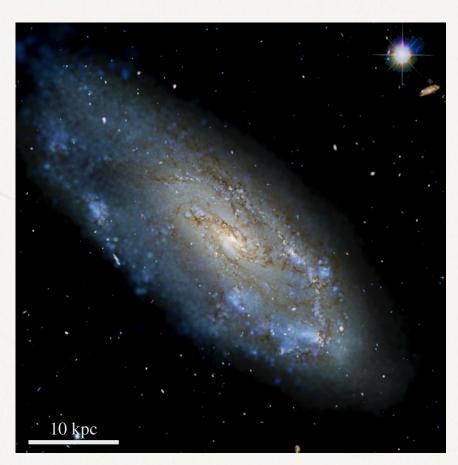
Opinion

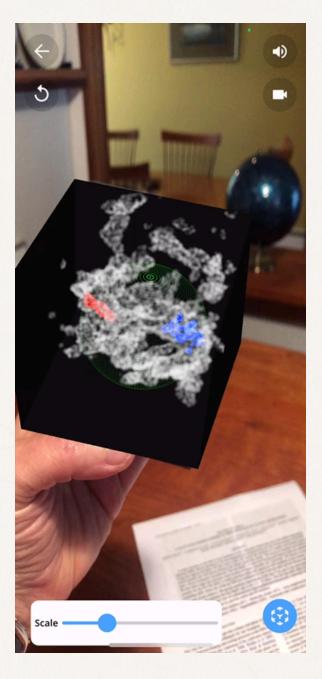
The New Universe

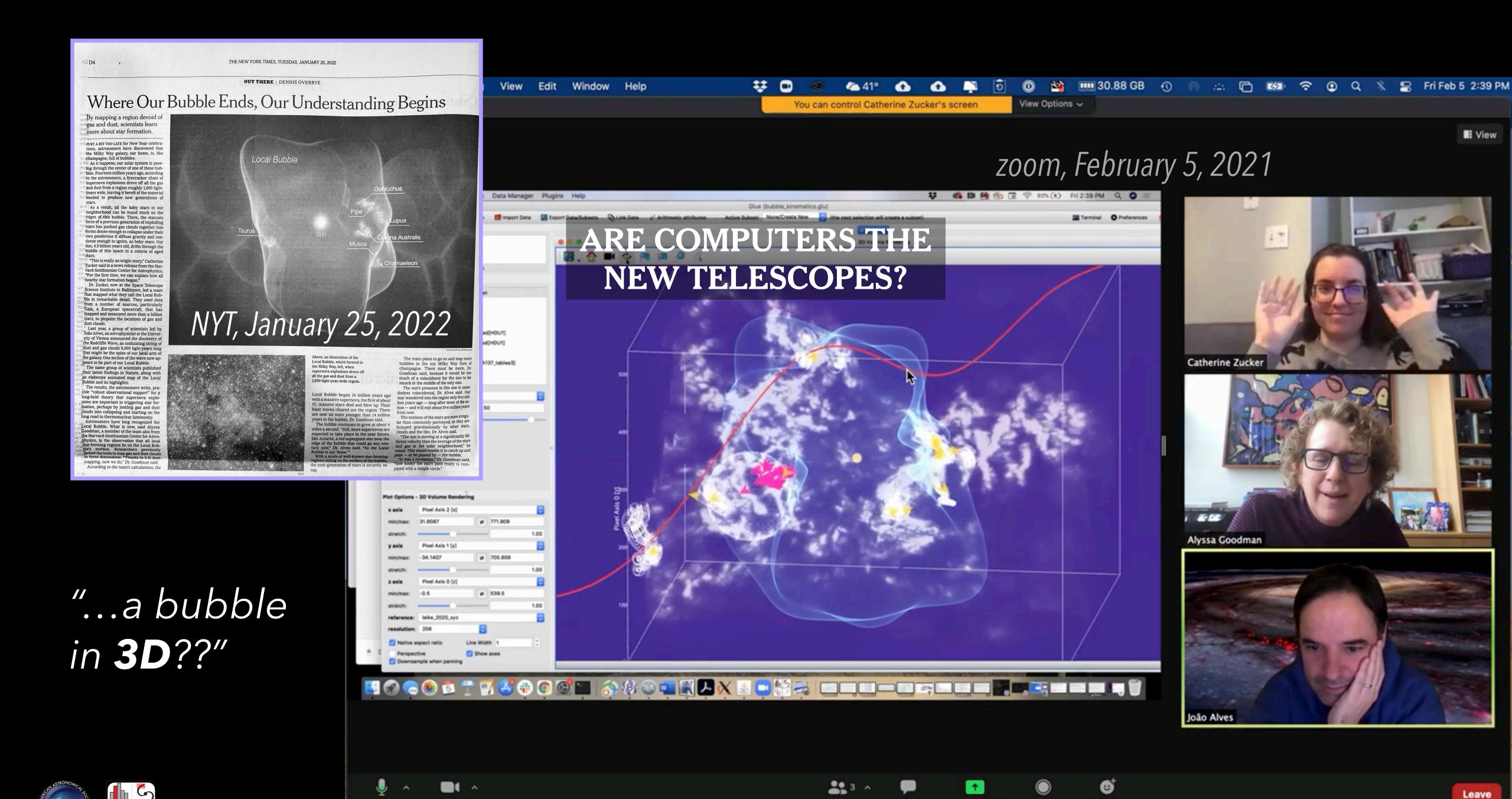
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Chat Share Screen Record

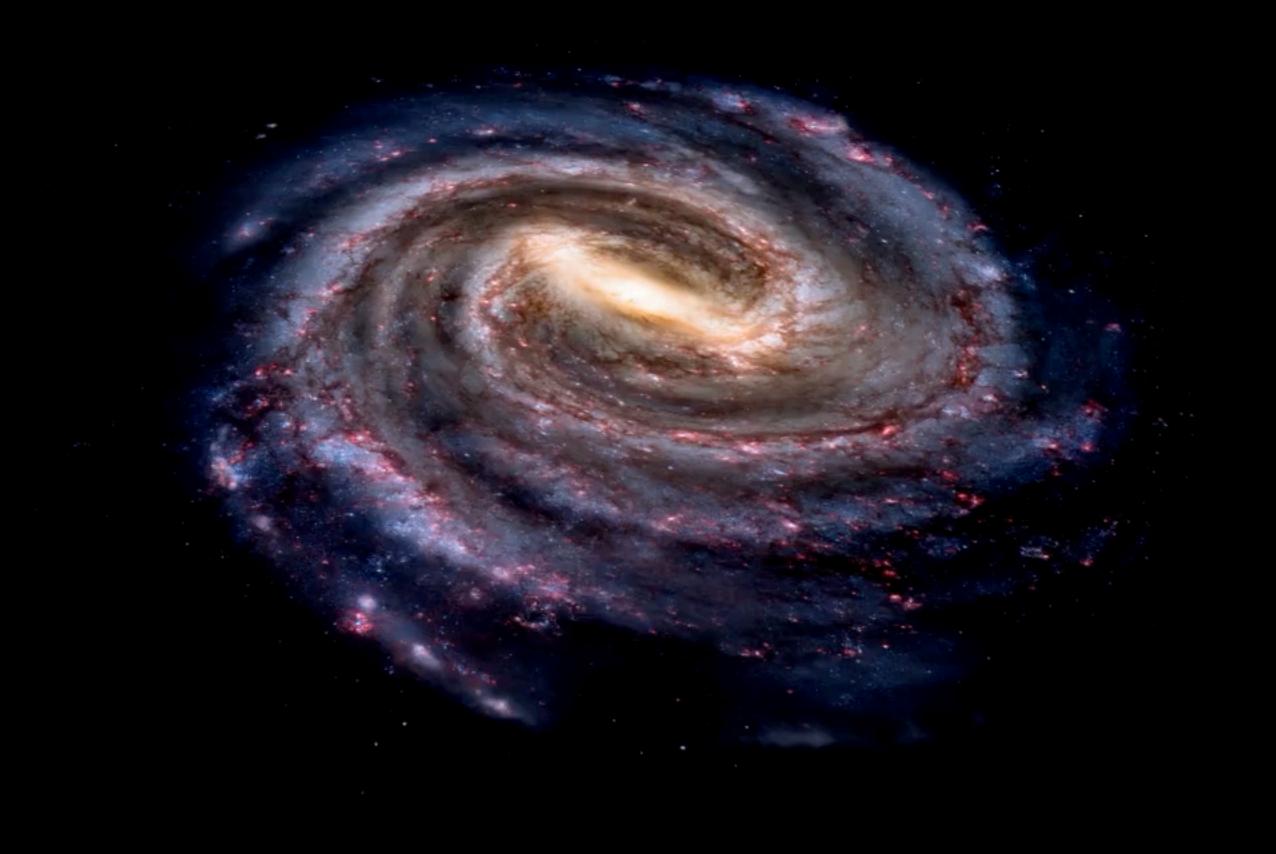
Participants

Reactions

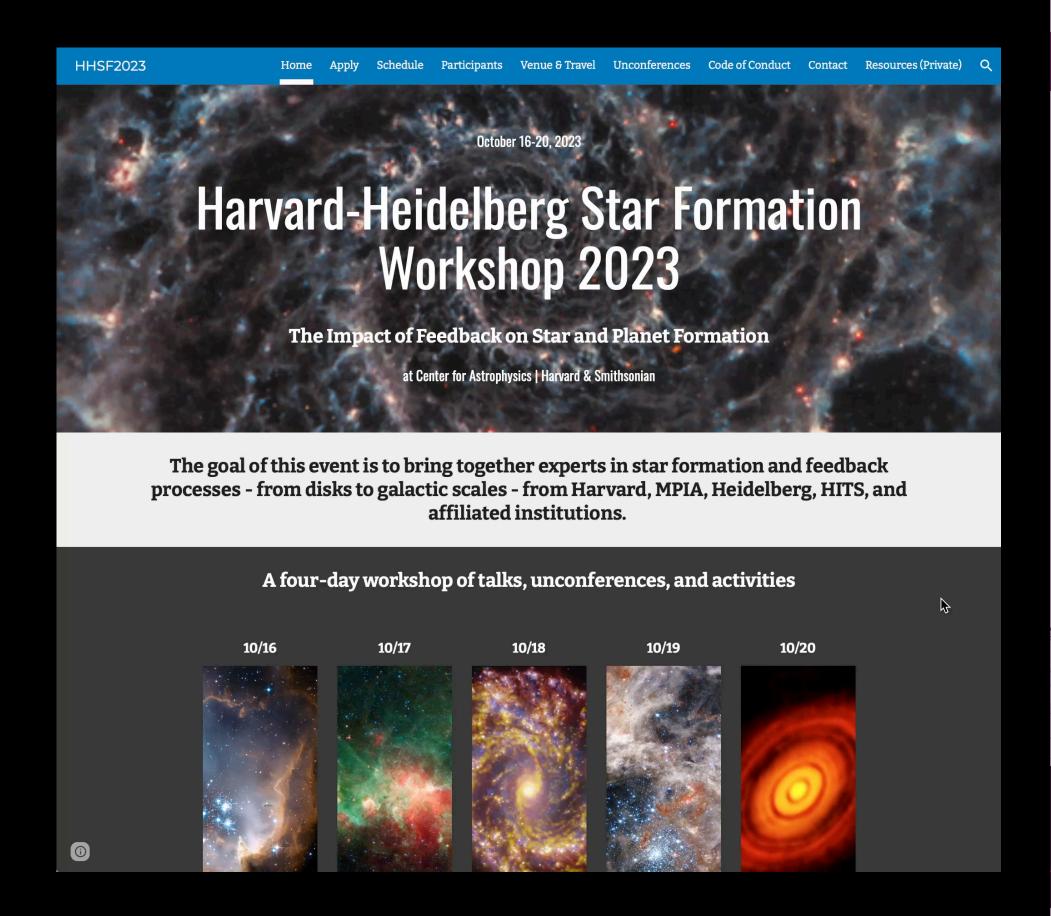
Stop Video

A Bubbly Origin

for Stars Around the Sun







J WAY JU. org

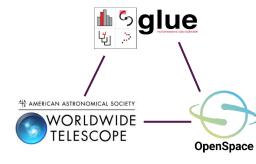


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enabling studies of how galaxies turn assembling data as a community, using modern, open-source practices



linking position and motion across dimensions, using the plug-in architecture enabled by glue (plug-ins include WorldWide Telescope, OpenSpace & more)







making data accessible online for decades

Lead: Alyssa Goodman, CfA

gas into stars, using...









topology, positions & motions of (long) features



tracers of feedback & magnetic fields



details on star-forming regions...and more!

Lead: Catherine Zucker, CfA

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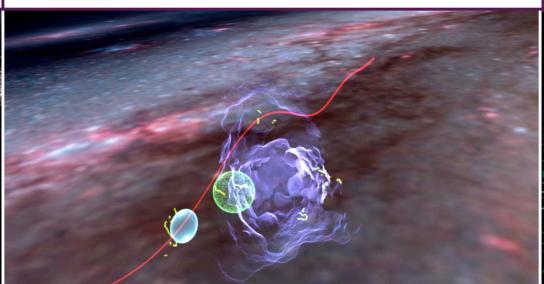
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"Cosmic Data Stories" teach data science using astronomical data & tools

Lead: Jackie Faherty, AMNH







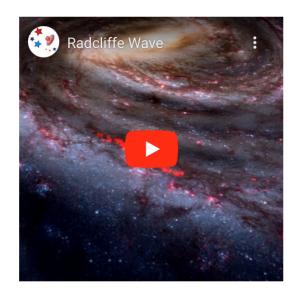
TEAM: Harvard/Smithsonian CfA (Jonathan Carifio, Alyssa Goodman, Ralf Konietzka, Theo O'Neill, Patricia Udomprasert, Catherine Zucker), AMNH (Brian Abbott, Micah Acinapura, Carter Emmart, Jackie Faherty); Linkoping University (Alex Bock) ; University of Vienna (Joao Alves, Sebastian Ratzenbock); glue solutions, inc./Aperio (Thomas Robitaille); University of Wisconsin, Whitewater (Bob Benjamin), STScI/Johns Hopkins (Josh Peek), Max Planck IfA (Gordian Edenhofer); Northeastern University (Michelle Borkin); and YOU

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THE SCIENCE STORY

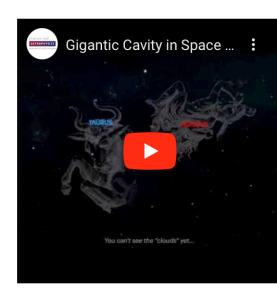
Several recent journal articles and data explorations by the scientists in our collaboration have motivated this project. The articles used a wide variety of data types and techniques to uncover several new 3D structures—some whose motions are also measured—in the Sun's neighborhood of the Milky Way, including The Radcliffe Wave, The Perseus-Taurus Supershell, the arrangement of star-forming regions on the surface of The Local Bubble around the Sun, new cavities created by star formation near the Orion nebula, and new compilations of young stars' properties, motions and positions. The combination of data and visualization from these and several related projects will be summarized in a ~2023 article entitled the "Charting the Nearby Milky Way in 3D," aimed at a "Scientific American" level audience. Data and interactive visualizations included in that initial article from our collaboration will be offered and archived publicly online, using the MAST, Dataverse, and milkyway3d.org mechanisms described in "The Data Story" above.



The Radcliffe Wave

The Radcliffe Wave is a gigantic structure that defines the shape of the Sun's local neighborhood in the Milky Way Galaxy. Its existence was first presented officially in a paper published in Nature on January 7, 2020. Its website offers scientists, educators, and the interested public much more information about the "RadWave," as we like to call it. Please use this page to find publications and talks, visuals (images,

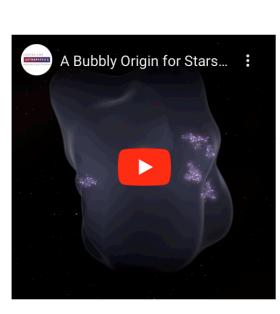
interactives, and videos), history, team info, software,



The Perseus-Taurus Supershell

Astronomers analyzing 3D maps of the shapes and sizes of nearby molecular clouds have discovered a gigantic cavity in space.

The sphere-shaped void, described in the Astrophysical Journal Letters, spans about 150 parsecs — nearly 500 light years — and is located on the sky among the constellations Perseus and Taurus. The research team believes the cavity was formed by ancient supernovae that went off some 10 million



The Local Bubble

The discovery that the 1000-light-year-wide "Local Bubble" surrounding the Sun and Earth is responsible for the formation of all nearby, young stars was first presented in a paper published in *Nature* on January 12, 2022. Please use this page to find news, publications and talks, visuals (images, interactives, and videos), team info, and data.

MILHY WAY 3D. org

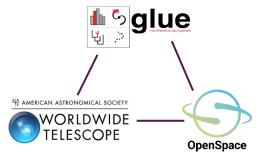


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enabling studies of how galaxies turn

gas into stars, using...

stellar positions, clusters, motions







topology, positions & motions of (long) features



tracers of feedback & magnetic fields



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EDULATION B OUTREACH

connecting real research data, software, and science to learners





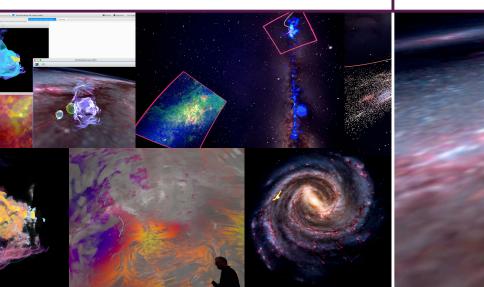


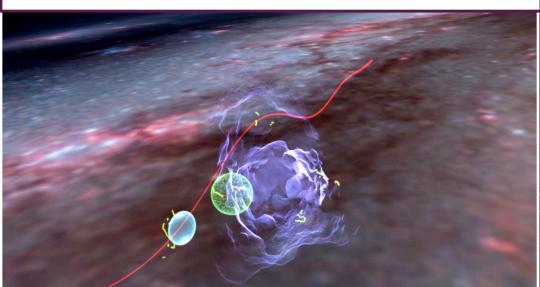
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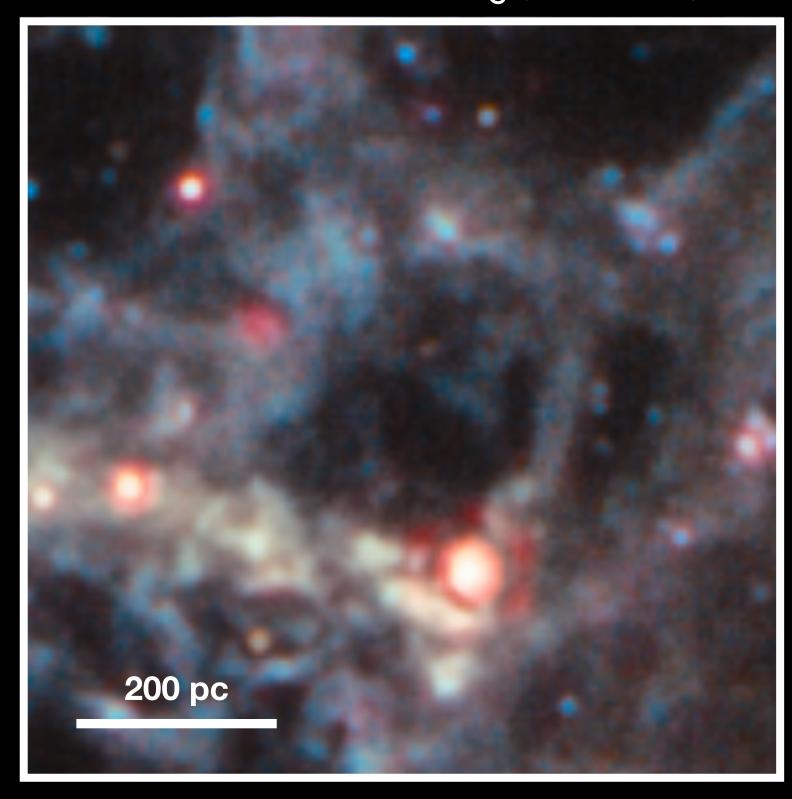


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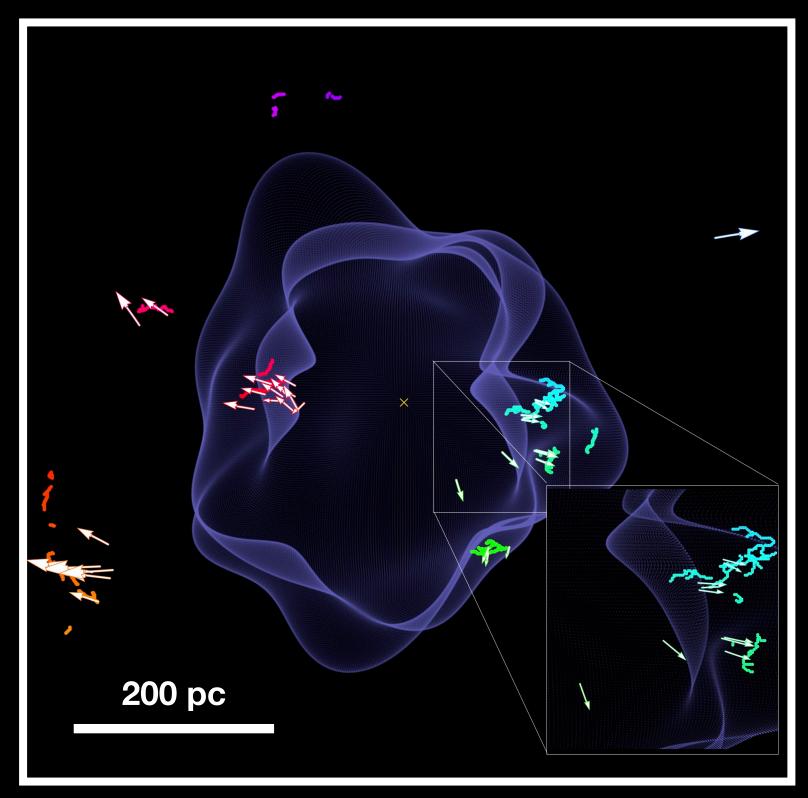
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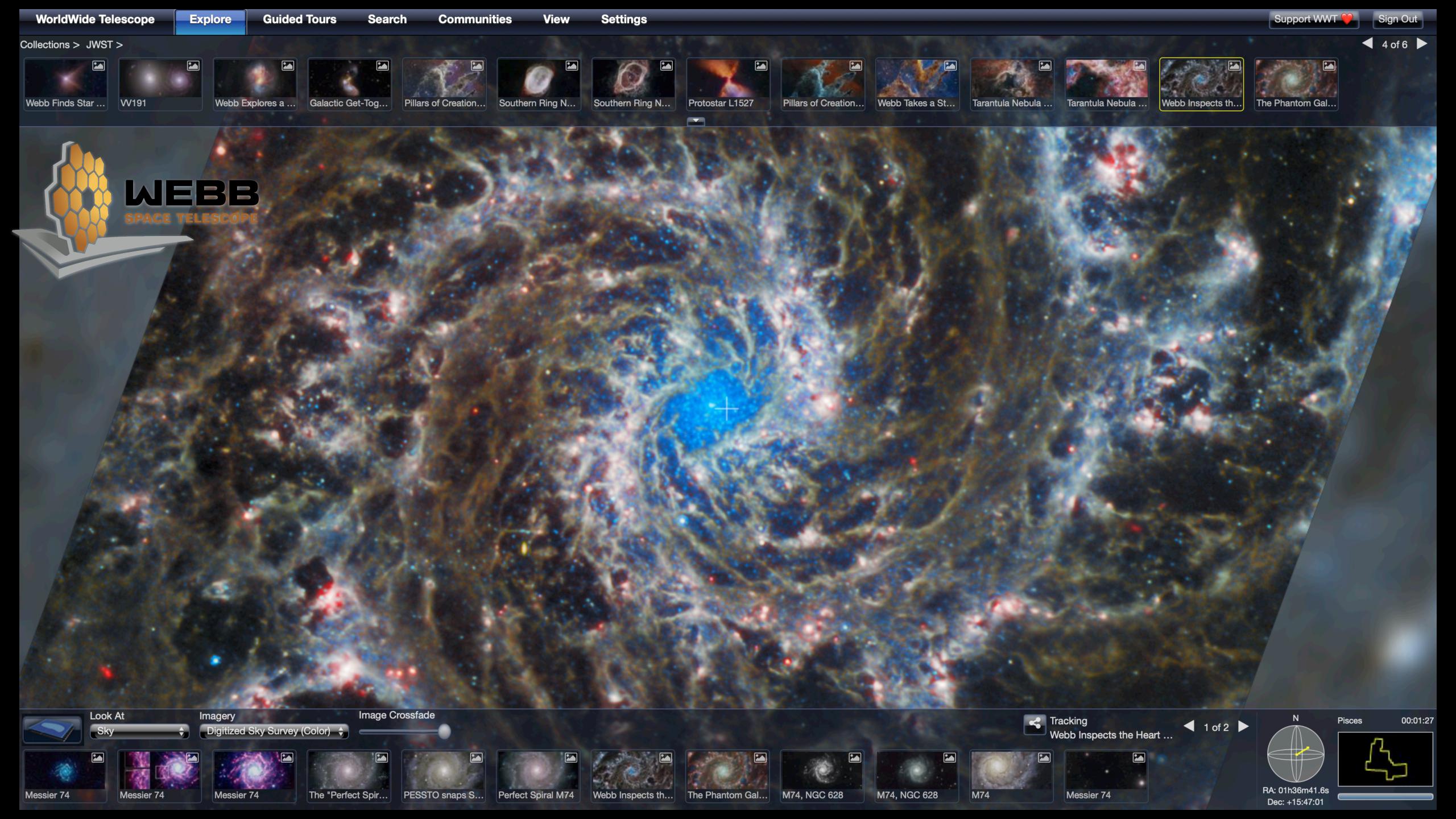


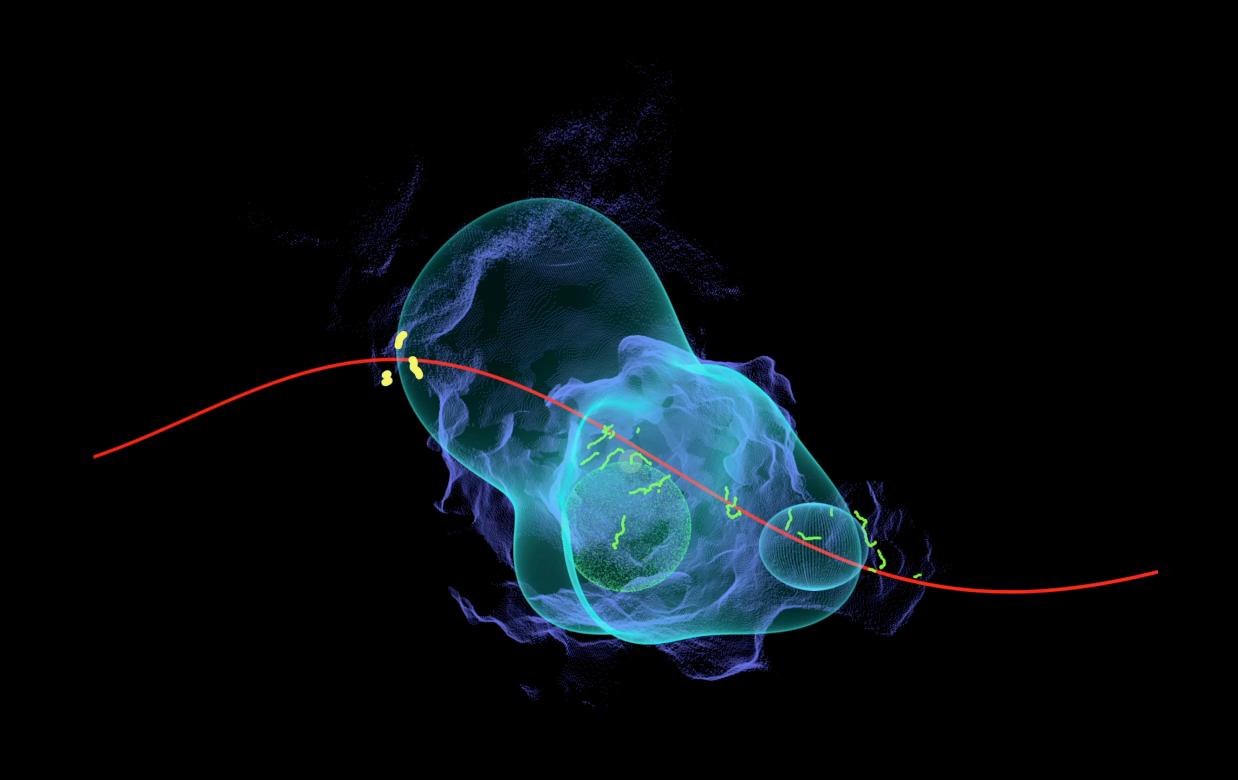
Local Bubble Analog (NGC 628)

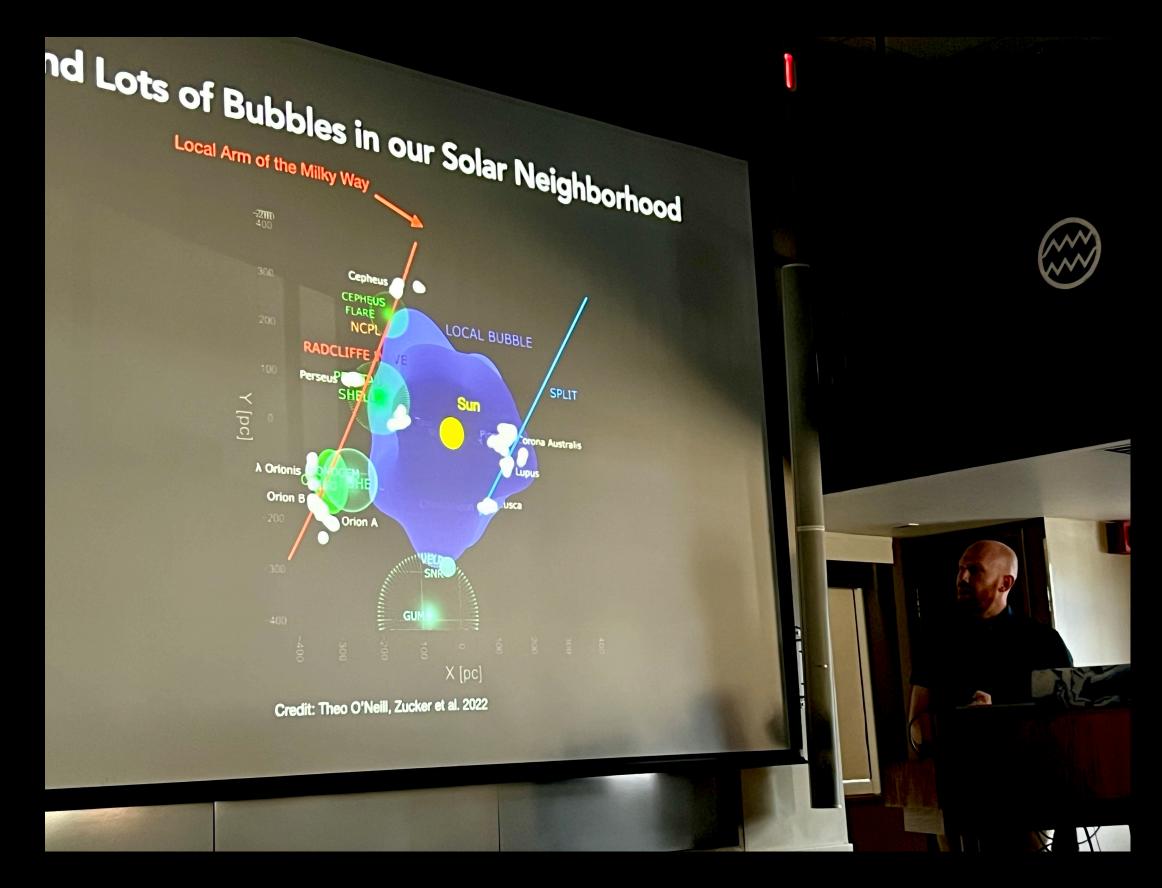


Local Bubble









Bubbles & the Radcliffe Wave in "glue"

Talk by CfA graduate student Michael Foley

"ARE COMPUTERS THE NEW TELESCOPES?"

HOW= 3D dust mapping*



+ Gaia*



+ glue*

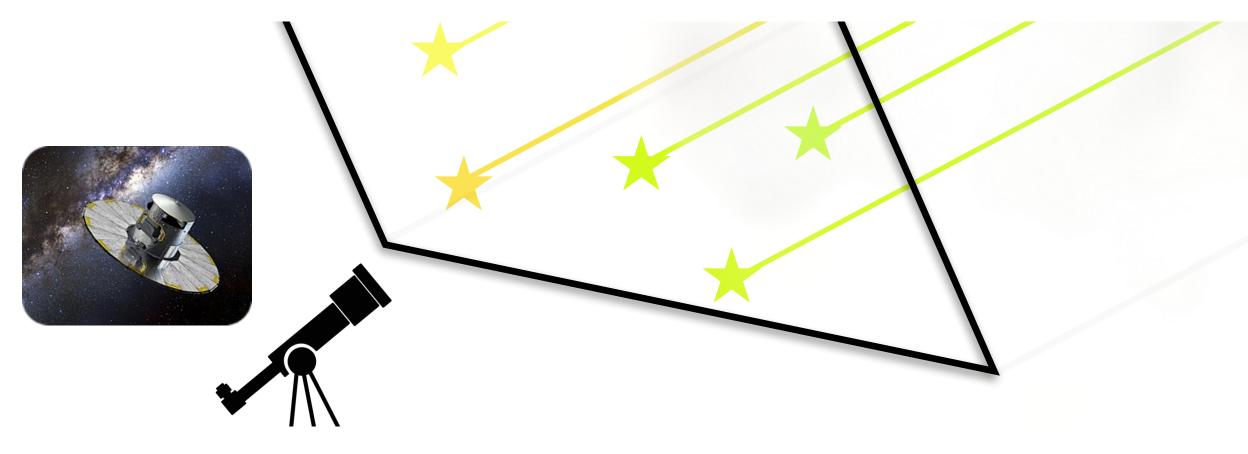


+ WorldWide Telescope

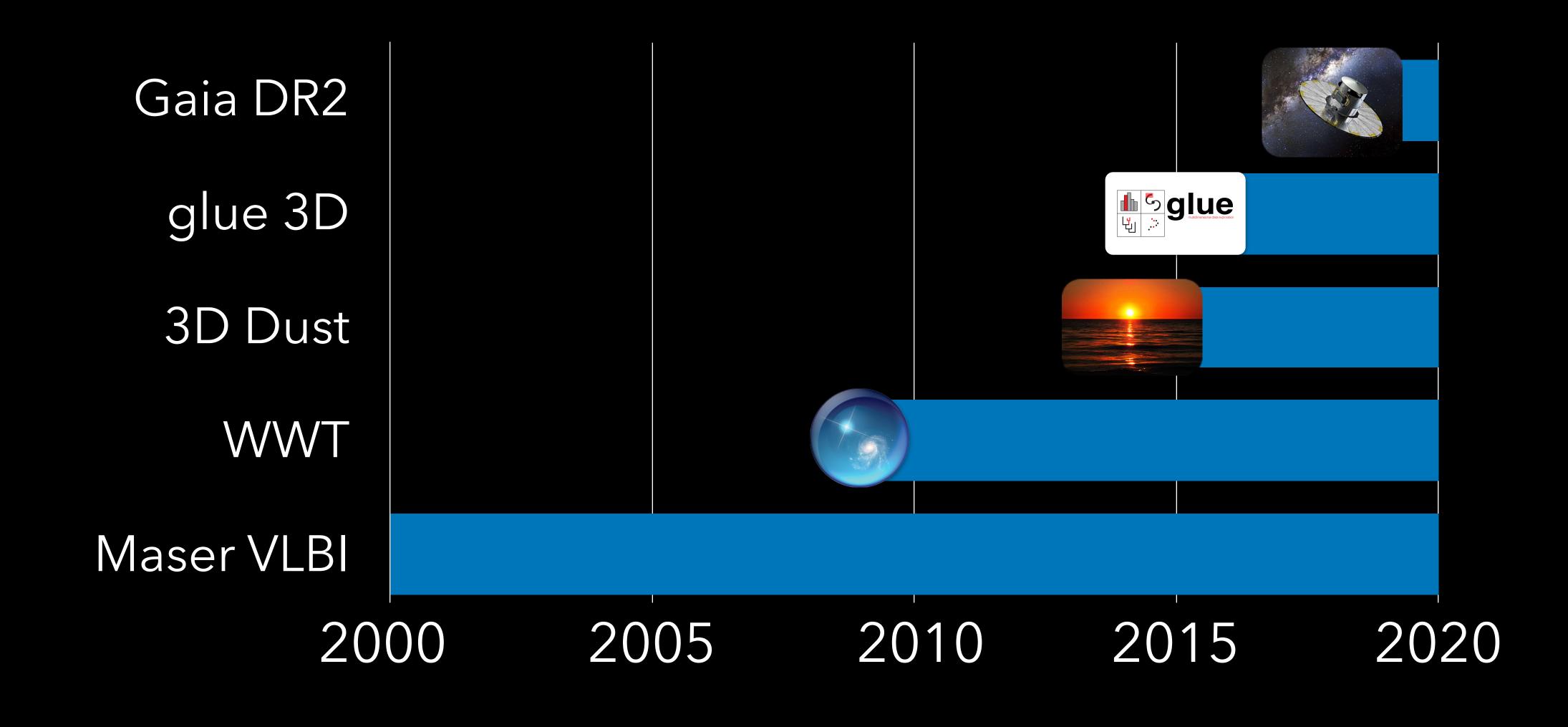




Can infer matter's distance from dust's effects on stars.



WHY COULDN'T WE DO THIS ALL A LONG TIME AGO?





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CALILEI



Aristotle, mon ami, above and below, it's all matter We're gonna need better math.



This apple.
That moon.
(ould they feel the same pull?

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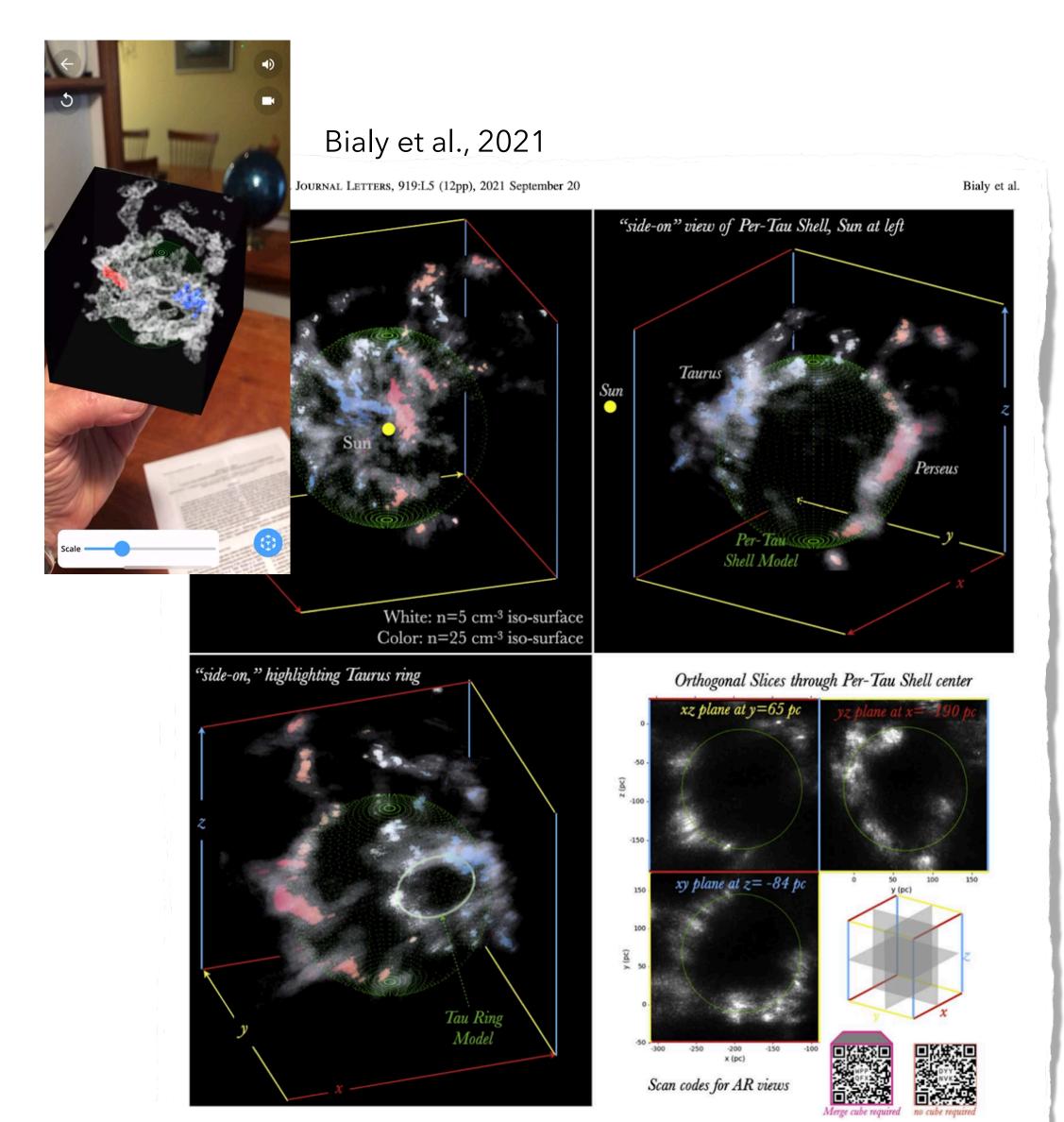


Figure 2. 3D views of the Per-Tau shell (for an interactive version⁸ of this figure click here ⁹; see Figure 5 for more static visualizations). Plotted are density isosurfaces at levels n = 5 cm⁻³ (gray) and n = 25 cm⁻³ (color), overlaid with our spherical-shell model, radius $R_s = 78$ pc, distance from the Sun d = 218 pc. The n = 25 cm⁻³ surfaces are colored by distance from the Sun (blue-to-red). Top-left panel: view from the Sun (compare with Figure 1). Top-right panel: a side view of the region. Perseus and Taurus and their diffuse envelopes are arranged on two opposing sides of the Per-Tau shell. Bottom-left panel: another side view emphasizing the Tau Ring. The ellipse is the Tau Ring model (Appendix B). Bottom-right panel: 2D density slices along the xy, xz, yz planes. All planes intersect at shell's center. In all panels xyz are the Heliocentric Cartesian Galactic Coordinates.

- 2. Tau Ring: in a sky projection the Tau Ring is seen almost edge-on. The near side of the Tau Ring connects with the main body of Taurus at $d \approx 150$ pc, whereas the farthest
- The Fictitious Connection: A filament seems to connect Taurus to Perseus. This connection is only a coincidental projection effect, where in actuality the filament is located

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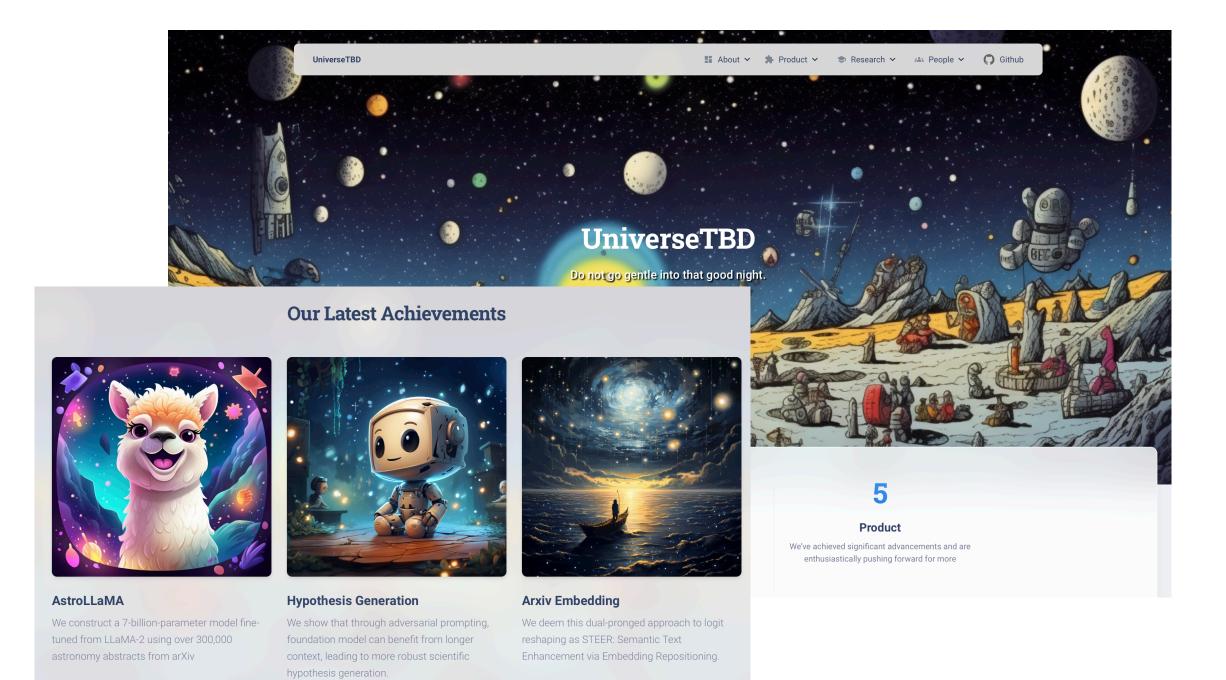
Developing Artificial Intelligence to Solve the Mysteries of the Universe

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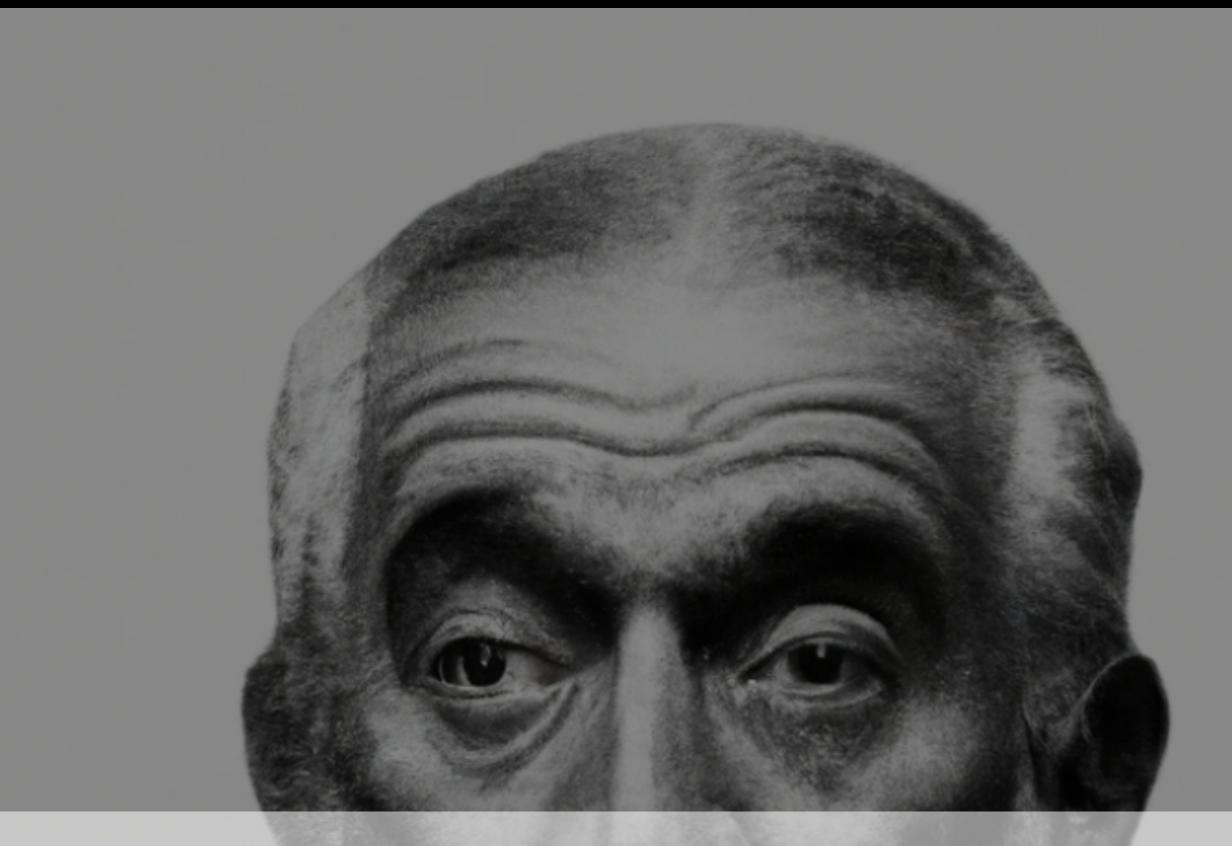
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Analysis: How AI is helping astronomers study the universe

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AI-assisted superresolution cosmological simulations

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Significance

Cosmological simulations are indispensable for understanding our Universe, from the creation of the cosmic web to the formation of galaxies and their central black holes. This vast dynamic range incurs large computational costs, demanding sacrifice of either resolution or size and often both. We build a deep neural network to enhance low-resolution dark-matter simulations, generating superresolution realizations that agree remarkably well with authentic high-resolution counterparts on their statistical properties and are orders-of-magnitude faster. It readily applies to larger volumes and generalizes to rare objects not present in the training data. Our study shows that deep learning and cosmological simulations can be a powerful combination to model the structure formation of our Universe over its full dynamic range.

